

Dynamical masses and mass-to-light ratios of resolved massive star clusters – II. Results for 26 star clusters in the Magellanic Clouds

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ABSTRACT

We present spectroscopy of individual stars in 26 Magellanic Cloud (MC) star clusters with the aim of estimating dynamical masses and V -band mass-to-light (M/L_V) ratios over a wide range in age and metallicity. We obtained 3137 high-resolution stellar spectra with M2FS on the *Magellan/Clay* Telescope. Combined with 239 published spectroscopic results of comparable quality, we produced a final sample of 2787 stars with good quality spectra for kinematic analysis in the target clusters. Line-of-sight velocities measured from these spectra and stellar positions within each cluster were used in a customized expectation-maximization (EM) technique to estimate cluster membership probabilities. Using appropriate cluster structural parameters and corresponding single-mass dynamical models, this technique ultimately provides self-consistent total mass and M/L_V estimates for each cluster. Mean metallicities for the clusters were also obtained and tied to a scale based on calcium IR triplet metallicities. We present trends of the cluster M/L_V values with cluster age, mass, and metallicity, and find that our results run about 40 per cent on average lower than the predictions of a set of simple stellar population (SSP) models. Modified SSP models that account for internal and external dynamical effects greatly improve agreement with our results, as can models that adopt a strongly bottom-light IMF. To the extent that dynamical evolution must occur, a modified IMF is not required to match data and models. In contrast, a bottom-heavy IMF is ruled out for our cluster sample as this would lead to higher predicted M/L_V values, significantly increasing the discrepancy with our observations.

Key words: techniques: radial velocities – techniques: spectroscopic – stars: abundances – stars: kinematics and dynamics – Magellanic Clouds – galaxies: star clusters: general.

1 INTRODUCTION

The baryonic mass-to-light (M/L) ratios of galaxies rely crucially on the detailed star-formation and chemical-enrichment histories of the multiple stellar populations that can exist in such stellar systems. As such, M/L ratios offer a convenient way to estimate the baryonic masses of galaxies from photometric observations, and numerous models to do so have been developed over the years (e.g. Bruzual & Charlot 2003; Maraston 2005; Conroy, Gunn & White 2009; Conroy & Gunn 2010; Vazdekis et al. 2010). The application of such models to data from large-scale photometric and spectroscopic surveys more-or-less directly determines much of what we know about how galaxies evolve with redshift (e.g. Bell et al. 2003; Kauffmann et al. 2003; Blanton & Roweis 2007; Tojeiro et al. 2009; Chen et al. 2012; Maraston et al. 2013), highlighting the need to test these models wherever possible.

To be useful, such tests should involve independent estimates of the masses of stellar systems from kinematic observations along with high-quality luminosity measurements. Moreover, an ideal test should involve systems with simple evolutionary histories so that M/L ratio can be tied to a specific age and metallicity (or, a small range in both). Although masses and luminosities can certainly be measured for galaxies, their complex evolutionary histories make them ill-suited to test model predictions of M/L ratios.

In contrast, star clusters provide a better alternative to test M/L predictions given their comparatively simple dynamical states and generally small internal variations of age and metallicity in a given system. Such tests boil down to using kinematic data to derive cluster masses and photometric measurements to estimate total luminosities. From both measurements, M/L ratios can be estimated in as purely empirical a manner as possible. Note that the questions such tests address go beyond determining which set of models is ‘best’, but can also explore to what extent – and, if so, why – M/L ratios vary among stellar systems of similar age and metallicity.

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From an observational standpoint, two spectroscopic methods can be used to determine the dynamical masses of star clusters: (1) measuring the central radial velocity dispersions from integrated-light spectroscopy (e.g. Illingworth 1976; Mandushev, Staneva & Spasova 1991; Zaritsky et al. 2012, 2013, 2014), and (2) collecting radial velocities of individual member stars (e.g. Gunn & Griffin 1979; Meylan & Mayor 1986; Lupton, Gunn & Griffin 1987; Lupton et al. 1989; Mateo, Welch & Fischer 1991; Fischer et al. 1992a; Fischer, Welch & Mateo 1992b, 1993; Suntzeff et al. 1992). In recent years, the second approach has become significantly more practical with the development of wide-field multi-object spectrographs (MOSs; e.g. Lane et al. 2010; Mackey et al. 2013; Kimmig et al. 2015; Song et al. 2019) and comparatively wide-field integral field units (IFUs; e.g. Kamann et al. 2016, 2018a,b). In addition to spectroscopy, proper motion data are also available for cluster mass determination – though mostly for the globular clusters in the Milky Way – via the *Hubble Space Telescope* (*HST*) observations (e.g. Bellini et al. 2014; Watkins et al. 2015) and the *Gaia* Data Release 2 (*Gaia* Collaboration 2018, hereafter *Gaia* DR2) astrometry (e.g. Baumgardt et al. 2019). In contrast to the mass estimates, luminosities of star clusters are determined almost exclusively using calibrated surface brightness/density profiles (e.g. McLaughlin & van der Marel 2005; Song et al. 2019). Together, masses and luminosities measured in these ways can be combined to produce empirical M/L estimates (e.g. McLaughlin & van der Marel 2005; Strader et al. 2009; Strader, Caldwell & Seth 2011; Kimmig et al. 2015; Baumgardt & Hilker 2018).

Simple stellar population (SSP) models can predict the M/L ratios of star clusters as a function of age and metallicity. Numerous studies have compared predictions from such models with empirical M/L ratios, mostly for old globular clusters either in the Milky Way (e.g. Pryor & Meylan 1993; McLaughlin & van der Marel 2005; Kimmig et al. 2015; Baumgardt 2017; Baumgardt & Hilker 2018; Dalglish et al. 2020) or Local Group (LG) galaxies (e.g. Larsen et al. 2002; Strader et al. 2009, 2011). Some of these studies have revealed variations of M/L ratios with respect to metallicity and cluster masses that do not appear to conform to model expectations (e.g. Strader et al. 2009; Strader et al. 2011; Kimmig et al. 2015; Baumgardt 2017; Dalglish et al. 2020).

Much less common are high-precision M/L measurements for young and intermediate-age clusters. Such systems can, in principle, greatly expand the parameter space in age and metallicity over which SSP models can be tested. At a distance of 50–60 kpc, the populous clusters of the Magellanic Clouds (MCs) provide excellent laboratories to broadly test SSP models in just this manner. These clusters are compact enough in the sky for both integrated-light spectroscopy and photometry, but also close enough to allow spectroscopy of individual stellar members in sufficient numbers to produce good-quality statistical samples. Further, the fact that both integrated-light and individual-star methods can be applied effectively to many MC clusters makes them particularly useful test cases to understand the relative systematics that may arise from each technique. Indeed, both approaches have been employed in past studies of MC clusters. Zaritsky et al. (2012, 2013, 2014) obtained integrated-light spectroscopic observations to measure M/L ratios for a sample of 17 clusters younger than 7 Gyr in the Large and Small Magellanic Clouds (LMC and SMC, respectively). To date, high-precision individual-star spectroscopic studies have been carried out for six additional MC clusters (Fischer et al. 1992a; Fischer et al. 1992b, 1993; Mackey et al. 2013; Kamann et al. 2018b; Song et al. 2019; Patrick et al. 2020).

In an earlier paper (Song et al. 2019, hereafter S19), we described the basic methodology of our survey and the analysis of the data as

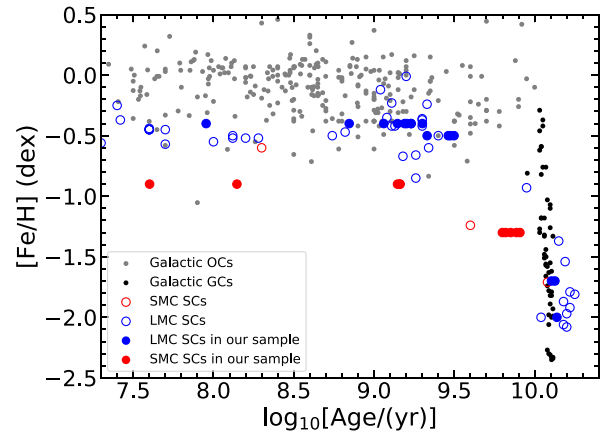


Figure 1. A representative plot of metallicity versus $\log(\text{Age})$ for star clusters in the Milky Way and the Magellanic Clouds. The filled circles denote the 16 LMC (blue) and 10 SMC (red) clusters studied in this paper, respectively. The metallicities and ages of these clusters correspond to the values listed in Table 1. The open circles correspond to other clusters in the LMC (blue) and SMC (red) from the collected catalogue of Pessev et al. (2006, 2008). For the Galactic clusters, we present the globular clusters listed in VandenBerg et al. (2013) (black dots) and the Galactic open clusters listed in Dias et al. (2002) (grey dots).

applied to the SMC cluster NGC 419 and the LMC cluster NGC 1846. Both clusters had been subjects of recent high-quality multi-object spectroscopic studies (Kamann et al. 2018b; Mackey et al. 2013) and so represented apt test cases to compare with independent empirical dynamical analyses. In general, we found acceptable agreement with past work once differences in kinematic precision were taken into account.

This paper updates the methodology from S19 as applied to our full kinematic sample consisting of 26 MC clusters with high-quality kinematic data (10 in the SMC, 16 in the LMC). These clusters were chosen to span the range from ~ 100 Myr to ~ 13 Gyr in age, and from -2.0 to -0.4 in $[\text{Fe}/\text{H}]$ (see Fig. 1 and Table 1) in order to provide the most leverage on our tests of V -band M/L (M/L_V) predictions from populations models. As in S19, this study employs spectroscopic observations obtained using the Michigan/Magellan Fiber System (M2FS) from which we measure kinematics and metallicities of samples of individual stars associated with the clusters in our sample. We derive dynamical masses and M/L_V ratios of all 26 clusters, along with independent spectroscopic estimates of the mean metallicities of the clusters.

This paper is organized as follows. The technical details of our study are described in Sections 2–4. Specifically, in Section 2 we describe our target selection, cluster centre determination, observational and data reduction procedures adopted for all clusters in our survey. Section 3 describes improvements of the Bayesian spectral fitting method we introduced in S19 to derive velocity and physical parameters from M2FS spectra. In Section 4, we obtain the final stellar sample for cluster kinematic analysis and assign a cluster membership probability to each star. In Section 5, we report the dynamical and chemical results for all clusters in our sample. We compare these results critically with those from previous studies. The key scientific results of our analysis are provided Section 6, where we discuss the trends of our M/L results with respect to various physical cluster parameters. We also compare these results with predictions from standard SSP models and with models that attempt to account for internal and external dynamical processes that can alter cluster

Table 1. General properties of star clusters in our sample.

Galaxy	Cluster	V_{ap} (mag)	Aper. ^a (arcsec)	Ref. ^b	Age (Gyr)	$(m - M)_0$ (mag)	A_V (mag)	Ref. ^b	[Fe/H] ^c (dex)
(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)	(9)	(10)
SMC	Kron 3	11.41 ± 0.09	50.0	1	7.10 ± 0.70	18.80 ± 0.04	0.08 ± 0.03	7	−1.3
SMC	Lindsay 1	13.32 ± 0.05	31.0	2	7.70 ± 0.70	18.69 ± 0.04	0.15 ± 0.03	7	−1.3
SMC	NGC 152	12.33 ± 0.05	50.0	1	1.40 ± 0.20	18.93 ± 0.04	0.16 ± 0.04	8	−0.9
SMC	NGC 330	9.60 ± 0.01	31.0	2	0.04 ± 0.00	18.80 ± 0.04	0.36 ± 0.04	9	−0.9
SMC	NGC 339	12.12 ± 0.04	50.0	1	6.30 ± 0.50	18.75 ± 0.08	0.19 ± 0.03	7	−1.3
SMC	NGC 361	12.24 ± 0.01	31.0	2	8.10 ± 1.20	18.49 ± 0.04	0.40 ± 0.04	8	−1.3
SMC	NGC 411	11.81 ± 0.07	50.0	1	1.45 ± 0.05	18.82 ± 0.03	0.16 ± 0.02	10	−0.9
SMC	NGC 416	11.42 ± 0.00	31.0	2	6.60 ± 0.80	18.76 ± 0.07	0.25 ± 0.03	7	−1.3
SMC	NGC 419	10.30 ± 0.16	50.0	1	1.45 ± 0.05	18.85 ± 0.03	0.15 ± 0.02	10	−0.9
SMC	NGC 458	11.73 ± 0.03	31.0	2	0.14 ± 0.03	19.11 ± 0.20	0.12 ± 0.06	11	−0.9
LMC	Hodge 4	13.33 ± 0.02	19.0	3	2.14 ± 0.00	18.37 ± 0.03	0.12 ± 0.04	12	−0.5
LMC	NGC 1466	11.59 ± 0.03	30.0	4	13.38 ± 2.00	18.66 ± 0.03	0.16 ± 0.04	13	−1.7
LMC	NGC 1751	11.67 ± 0.13	50.0	1	1.40 ± 0.05	18.50 ± 0.03	0.38 ± 0.02	10	−0.4
LMC	NGC 1783	10.39 ± 0.03	50.0	1	1.70 ± 0.05	18.49 ± 0.03	0.00 ± 0.02	10	−0.4
LMC	NGC 1806	11.00 ± 0.05	50.0	1	1.60 ± 0.05	18.50 ± 0.03	0.05 ± 0.03	10	−0.4
LMC	NGC 1831	11.18 ± 0.02	30.0	4	0.70 ± 0.10	18.23 ± 0.09	0.03 ± 0.06	14	−0.4
LMC	NGC 1841	11.43 ± 0.02	93.5	3	13.77 ± 1.70	18.34 ± 0.04	0.35 ± 0.04	13	−2.0
LMC	NGC 1846	10.68 ± 0.20	50.0	1	1.70 ± 0.05	18.42 ± 0.03	0.07 ± 0.02	10	−0.4
LMC	NGC 1850	9.57 ± 0.20	25.0	3	0.09 ± 0.05	18.45 ± 0.03	0.37 ± 0.02	15	−0.4
LMC	NGC 1978	10.20 ± 0.02	50.0	1	2.00 ± 0.00	18.55 ± 0.04	0.16 ± 0.04	16	−0.4
LMC	NGC 2121	12.37 ± 0.01	31.0	5	2.90 ± 0.50	18.24 ± 0.04	0.22 ± 0.06	14	−0.5
LMC	NGC 2155	12.59 ± 0.48	50.0	1	3.00 ± 0.25	18.32 ± 0.04	0.06 ± 0.03	14	−0.5
LMC	NGC 2203	11.29 ± 0.15	75.0	3	1.55 ± 0.05	18.37 ± 0.03	0.16 ± 0.02	10	−0.4
LMC	NGC 2209	13.15 ± 0.01	34.0	3	1.15 ± 0.05	18.37 ± 0.03	0.23 ± 0.02	17	−0.4
LMC	NGC 2257	12.62 ± 0.02	30.5	3	12.74 ± 2.00	18.25 ± 0.04	0.12 ± 0.04	13	−1.7
LMC	SL 663	22.13 ± 0.24 ^d	0.0	6	3.15 ± 0.40	18.32 ± 0.07	0.22 ± 0.06	14	−0.5

Notes. ^aAperture radius used for measuring the V_{ap} magnitude in column 3.

^bReferences: (1) Goudfrooij et al. (2006); (2) Alcaïno (1978); (3) Bica et al. (1996); (4) van den Bergh (1981); (5) Bernard (1975); (6) McLaughlin & van der Marel (2005); (7) Glatt et al. (2008); (8) Crowl et al. (2001); (9) Milone et al. (2018); (10) Goudfrooij et al. (2014); (11) Alcaïno et al. (2003); (12) Grocholski et al. (2007); (13) Wagner-Kaiser et al. (2017); (14) Kerber, Santiago & Brocato (2007); (15) Correnti et al. (2017); (16) Martocchia et al. (2018); (17) Correnti et al. (2014).

^cThe adopted [Fe/H] values were estimated from the age–metallicity relations for the LMC and SMC clusters, respectively. For the LMC clusters, we assumed that [Fe/H] = −0.4 if 0–2 Gyr, [Fe/H] = −0.5 if 2–4 Gyr, [Fe/H] = −1.7 for NGC 1466 and NGC 2257, and [Fe/H] = −2.0 for NGC 1841; while for the SMC clusters, we assumed that [Fe/H] = −0.9 if 0–4 Gyr, and [Fe/H] = −1.3 if 6–9 Gyr. These age–metallicity relations were averaged from multiple papers that fitted cluster CMDs with the Padova isochrones (Grocholski et al. 2007; Kerber et al. 2007; Milone et al. 2009; Goudfrooij et al. 2014; Milone et al. 2018 for LMC, and Crowl et al. 2001; Glatt et al. 2008; Goudfrooij et al. 2014; Milone et al. 2018 for SMC). The final cluster metallicities from this study are listed in Table 6.

^dSL 663 has no V-band aperture photometry in the literature. We have adopted the best-fitting V-band extinction-corrected central surface brightness from McLaughlin & van der Marel (2005). The value listed here is in unit of mag arcsec^{−2}.

M/L ratios. Section 7 summarizes our methodology and results, and provides a brief outline of our conclusions.

2 DATA OVERVIEW

2.1 Cluster candidates

We selected cluster candidates from catalogues of MC clusters with good-quality age and metallicity estimates and for which we could expect to obtain samples of a few dozen stellar members. Table 1 lists the general properties of the star clusters in our final sample and specifies all the literature sources used to select our spectroscopic targets. Columns 1 and 2 in the table specify the host galaxy and the most common names of each cluster. The V-band aperture magnitude is listed in column 3, with the corresponding aperture radius listed in column 4. Column 5 lists the sources for these photometric results. Columns 6 through 8 list age, distance modulus, and extinction values taken from the sources listed in column 9; these three parameters were estimated from the cited sources from comparisons of a given

cluster’s colour–magnitude diagram (CMD) with modern synthetic isochrones (see below). We gave preference to clusters with deep *HST* photometry, but if no *HST* data were available good-quality ground-based photometry was also used. Most of the studies listed in Table 1 used the isochrones of the Padova group (Girardi et al. 2000, 2002; Bressan et al. 2012). In a few cases, the sources listed in column 9 employed multiple sets of isochrones from different synthetic groups; for consistency, we adopted only their best-fitting results using Padova/PARSEC isochrones. Column 10 lists the adopted metallicities that estimated from the age–metallicity relations for the LMC and SMC clusters, respectively (see the corresponding table note for details). Fig. 1 plots the ages and metallicities from Table 1 of our sample, illustrating the wide range of these parameters sampled by the clusters in our study.

Table 2 lists the positions and structural parameters of the clusters in our sample. These parameters are essential for the background/sky subtraction processes (see Section 2.4.2) and to establish the central velocity dispersions of the clusters (see Section 4). The cluster positions listed in Table 2 (columns 3 and 4) were derived from

Table 2. Positions and structural parameters of star clusters in our sample.

Galaxy	Cluster	α_{J2000} (hh mm ss)	δ_{J2000} (dd mm ss)	c_{K62}^a	$r_{0,K62}$ (arcsec)	Ref. ^b	c_{K66}^a	$r_{0,K66}$ (arcsec)	Ref. ^b
(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)	(9)	(10)
SMC	Kron 3	00 24 45.98	-72 47 37.9	0.79 ^{+0.01} _{-0.01}	29.62 ^{+0.69} _{-0.69}	1	1.14 ^{+0.05} _{-0.04}	23.24 ^{+1.26} _{-1.31}	7
SMC	Lindsay 1	00 03 53.50	-73 28 15.0	0.55 ^{+0.02} _{-0.02}	66.76 ^{+2.50} _{-2.50}	1	0.60 ^{+0.02} _{-0.02}	75.40 ^{+2.82} _{-2.82}	...
SMC	NGC 152	00 32 56.76	-73 06 56.4	0.98 ^{+0.08} _{-0.09}	26.70 ^{+2.65} _{-1.77}	...	1.09 ^{+0.09} _{-0.10}	27.39 ^{+2.72} _{-1.82}	7
SMC	NGC 330	00 56 18.55	-72 27 45.1	1.35 ^{+0.09} _{-0.08}	8.34 ^{+0.74} _{-0.69}	...	1.41 ^{+0.09} _{-0.08}	8.55 ^{+0.76} _{-0.71}	7
SMC	NGC 339	00 57 47.62	-74 28 14.6	0.79 ^{+0.01} _{-0.01}	32.91 ^{+0.59} _{-0.59}	1	0.82 ^{+0.21} _{-0.17}	33.18 ^{+6.00} _{-4.34}	7
SMC	NGC 361	01 02 11.11	-71 36 25.3	0.88 ^{+0.24} _{-0.18}	26.55 ^{+4.90} _{-3.63}	...	0.99 ^{+0.27} _{-0.20}	27.27 ^{+5.03} _{-3.73}	7
SMC	NGC 411	01 07 55.64	-71 46 03.1	0.87 ^{+0.03} _{-0.03}	14.50 ^{+0.89} _{-0.89}	2	1.38 ^{+0.11} _{-0.11}	9.30 ^{+1.09} _{-0.99}	7
SMC	NGC 416	01 07 59.03	-72 21 20.5	0.81 ^{+0.02} _{-0.02}	12.56 ^{+0.48} _{-0.48}	1	0.89 ^{+0.08} _{-0.07}	12.05 ^{+0.86} _{-0.79}	7
SMC	NGC 419	01 08 17.26	-72 53 01.8	1.11 ^{+0.03} _{-0.03}	14.51 ^{+0.93} _{-0.93}	1, 2	1.20 ^{+0.03} _{-0.03}	14.90 ^{+0.95} _{-0.95}	...
SMC	NGC 458	01 14 52.94	-71 32 60.0	0.98 ^{+0.10} _{-0.11}	12.01 ^{+1.84} _{-1.36}	...	1.09 ^{+0.11} _{-0.12}	12.32 ^{+1.89} _{-1.40}	7
LMC	Hodge 4	05 32 25.64	-64 44 07.7	6.54 ^{+5.10} _{-1.25}	14.27 ^{+1.12} _{-0.61}	...	2.45 ^{+1.91} _{-0.47}	14.41 ^{+1.13} _{-0.62}	7
LMC	NGC 1466	03 44 32.71	-71 40 18.0	1.00 ^{+0.06} _{-0.05}	10.81 ^{+0.68} _{-0.65}	...	1.11 ^{+0.07} _{-0.06}	11.06 ^{+0.70} _{-0.66}	7
LMC	NGC 1751	04 54 12.91	-69 48 26.8	0.84 ^{+0.03} _{-0.03}	22.00 ^{+1.57} _{-1.57}	2, 3	0.95 ^{+0.03} _{-0.03}	22.64 ^{+1.61} _{-1.61}	...
LMC	NGC 1783	04 59 08.75	-65 59 15.6	0.96 ^{+0.02} _{-0.02}	37.70 ^{+1.76} _{-1.76}	2, 3	1.08 ^{+0.02} _{-0.02}	38.58 ^{+1.80} _{-1.80}	...
LMC	NGC 1806	05 02 11.86	-67 59 08.5	0.90 ^{+0.02} _{-0.02}	24.50 ^{+1.12} _{-1.12}	2, 3	1.01 ^{+0.02} _{-0.02}	25.15 ^{+1.15} _{-1.15}	...
LMC	NGC 1831	05 06 16.12	-64 55 06.9	1.08 ^{+0.05} _{-0.05}	18.24 ^{+0.90} _{-1.06}	...	1.18 ^{+0.05} _{-0.05}	18.65 ^{+0.92} _{-1.08}	7
LMC	NGC 1841	04 45 24.38	-83 59 53.3	0.53 ^{+0.26} _{-0.53}	42.76 ^{+77.76} _{-9.61}	...	0.57 ^{+0.28} _{-0.57}	49.14 ^{+89.36} _{-11.04}	7
LMC	NGC 1846	05 07 34.47	-67 27 37.8	0.79 ^{+0.03} _{-0.03}	26.00 ^{+1.59} _{-1.59}	2, 4	0.90 ^{+0.03} _{-0.03}	26.92 ^{+1.64} _{-1.64}	...
LMC	NGC 1850	05 08 46.32	-68 45 38.7	0.99 ^{+0.02} _{-0.02}	12.40 ^{+0.49} _{-0.49}	5	0.98 ^{+0.02} _{-0.02}	12.72 ^{+0.51} _{-0.51}	...
LMC	NGC 1978	05 28 44.73	-66 14 09.3	1.16 ^{+0.05} _{-0.05}	17.30 ^{+0.78} _{-0.78}	...	1.26 ^{+0.05} _{-0.05}	17.70 ^{+0.80} _{-0.80}	8
LMC	NGC 2121	05 48 12.74	-71 28 45.2	0.60 ^{+0.13} _{-0.13}	43.21 ^{+9.28} _{-5.90}	...	0.67 ^{+0.15} _{-0.15}	47.30 ^{+10.16} _{-6.46}	7
LMC	NGC 2155	05 58 31.98	-65 28 41.0	0.84 ^{+0.17} _{-0.14}	19.20 ^{+2.78} _{-1.98}	...	0.95 ^{+0.19} _{-0.16}	19.76 ^{+2.86} _{-2.04}	7
LMC	NGC 2203	06 04 42.17	-75 26 15.8	0.69 ^{+0.02} _{-0.02}	32.90 ^{+1.61} _{-1.61}	2	0.79 ^{+0.02} _{-0.02}	34.61 ^{+1.69} _{-1.69}	...
LMC	NGC 2209	06 08 36.06	-73 50 07.9	0.85 ^{+0.02} _{-0.02}	27.30 ^{+1.40} _{-1.40}	6	0.83 ^{+0.02} _{-0.02}	28.09 ^{+1.44} _{-1.44}	...
LMC	NGC 2257	06 30 12.62	-64 19 40.0	0.81 ^{+0.30} _{-0.21}	30.99 ^{+6.33} _{-3.88}	...	0.91 ^{+0.34} _{-0.24}	32.02 ^{+6.54} _{-4.01}	7
LMC	SL 663	05 42 28.20	-65 21 50.2	2.19 ^{+2.93} _{-1.14}	27.85 ^{+7.10} _{-1.57}	...	1.86 ^{+2.49} _{-0.97}	28.58 ^{+7.29} _{-1.61}	7

Notes. ^aConcentration parameter $c \equiv \log_{10} r_t / r_0$, where r_0 and r_t are the King radius and truncation radius, respectively.

^bReferences: (1) Glatt et al. (2009); (2) Goudfrooij et al. (2014); (3) Goudfrooij et al. (2011); (4) Goudfrooij et al. (2009); (5) Correnti et al. (2017); (6) Correnti et al. (2014); (7) McLaughlin & van der Marel (2005); (8) Fischer et al. (1992b).

the *Gaia* DR2. Details of how we determined these centres can be found in Appendix A.¹ For the cluster structural parameters, columns 5 and 6 give the concentration parameter and the King radius of an empirical number density profile by King (1962, hereafter the **K62** profile). Column 7 lists the references for these structural parameters. Columns 8 and 9 give the same parameters but for a dynamical model developed by King (1966, hereafter the **K66** model). In cases where only the **K62** profile or **K66** model is available from the literature, we transformed the structural parameters from one to the other, as described in S19. This approach ensured that the central surface brightness, core radius (defined as the radius at the half of the central surface brightness) and total luminosity agree for both profiles.

2.2 Target selection within cluster fields

For any given cluster, we selected a variety of specific types of targets for spectroscopic analysis. The primary science targets were

typically drawn from the red giant branch (RGB) of a cluster's CMD. These targets were prioritized according to their proximity to their respective cluster centre. Additional science targets beyond the formal tidal radii of clusters were also included to allow us to determine the kinematic and chemical distribution of the local field populations. For both science target selections, we identified apparently isolated stars as potential spectroscopic targets. We regarded a star to be isolated when the integrated flux of all other stars in the corresponding photometric catalogue within 1 arcsec of the star adds up to ≤ 20 per cent of the candidate star's flux.

The scientific targets were selected in the following manner. Before the *Gaia* DR2 became available on 2018 April, our target selection generally relied on archival *Hubble Space Telescope* (*HST*) images (see below), the Magellanic Clouds Photometric Survey (MCPS) *UBVI* catalogue (Zaritsky et al. 2002, 2004), and the Johnson *BV* photometry by Jeon et al. (2014). Some of these sources exhibited rather large astrometric offsets and/or scale errors, which made the observational data based on these sources less reliable. For all adopted *HST* images, short exposures with the F555W and F814W (or F450W when F814W was not available) filters were reduced using the WFPC2 or ACS modules from the DOLPHOT package (Dolphin 2000), and the resulting CMDs were mainly used to select

¹We have not tested the impact of centring errors on the cluster structural parameters. Hence the robustness of the central velocity dispersion measurements (see Section 5.1) relative to small shifts of the cluster centres is based on the assumption that the structural parameters are unchanged to such shifts.

Table 3. Observations.

Galaxy	Cluster	Obs. Date (UT)	Exp. Time (on target) (s)	Exp. Time (offset) (s)	N_{star}	Sources ^a	N_{sky}	Sky Sub. Method
(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)	(9)
SMC	Kron 3	2018-08-22	3 × 1200	2 × 300	120	1	12	A
SMC	Lindsay 1	2018-08-21	1 × 1200 + 1 × 1500 + 1 × 1200	1 × 600	120	1	12	A
SMC	NGC 152	2018-08-20	3 × 1200	2 × 300	120	1	12	A
SMC	NGC 330	2018-08-18	1 × 1200 + 2 × 1500	2 × 480	39	1	15	A
SMC	NGC 339	2016-09-12	5 × 2400	2 × 600	127	2, 3	3	A
SMC	NGC 361	2018-11-26	3 × 1200	3 × 600	120	1	12	A*
SMC	NGC 411	2018-08-15	3 × 1500	2 × 600	120	1	12	A
SMC	NGC 416	2018-08-21	3 × 1200	2 × 300	120	1	12	A
SMC	NGC 419	2017-09-21	4 × 2100	2 × 480	123	2, 3	5	A
SMC	NGC 458	2018-08-15	3 × 1200	2 × 300	82	1	16	A
LMC	Hodge 4	2018-12-02	2 × 1080 + 1 × 840	2 × 240	119	1	12	A
LMC	NGC 1466	2016-12-11	3 × 2400	0	46	4, 5	42	B
LMC	NGC 1751	2019-02-24	3 × 1730	2 × 450	118	1	12	A
LMC	NGC 1783	2018-02-16	1 × 1800 + 2 × 2000	0	114	2, 6	9	B
LMC	NGC 1806	2018-11-30	1 × 1200 + 2 × 960	2 × 240	120	1	12	A
LMC	NGC 1831	2018-12-01	1 × 1500 + 1 × 1200 + 1 × 1500	2 × 300	119	1	12	A
LMC	NGC 1841	2018-02-25	2 × 1200	0	112	4	12	B
LMC	NGC 1846	2018-02-21	4 × 1200	0	111	2, 6	13	B
LMC	NGC 1850	2018-02-19	3 × 1200	0	113	2, 7	13	B
LMC	NGC 1978	2017-03-03	3 × 1800	0	130	2, 8	1	B
LMC	NGC 1978	2017-11-10	3 × 2000	0	129	2, 8	0	–
LMC	NGC 2121	2018-02-23	3 × 1200	0	117	2, 7	8	B
LMC	NGC 2155	2018-12-08	3 × 1800	2 × 600	119	1	12	A
LMC	NGC 2203	2018-12-04	1 × 1200 + 1 × 1000 + 2 × 1800	4 × 600	119	1	12	A
LMC	NGC 2203	2019-03-06	1 × 1600 + 2 × 1200	2 × 600	118	1	12	A
LMC	NGC 2209	2018-12-05	1 × 1200 + 1 × 1800	1 × 480 + 1 × 600	119	1	11	A
LMC	NGC 2209	2019-03-01	3 × 1850	0	118	1	11	B
LMC	NGC 2257	2018-05-18	3 × 1800	0	116	4	8	B
LMC	SL 663	2019-02-28	3 × 1750	0	118	1	12	B

Notes. ^aSources for target selection: (1) GDR2 (Gaia Collaboration 2018); (2) MCPS (Zaritsky et al. 2002, 2004); (3) *HST* GO-10396 (PI: Gallagher); (4) Jeon et al. (2014); (5) *HST* GO-5897 (PI: Bolte); (6) *HST* GO-10595 (PI: Goudfrooij); (7) *HST* SNAP-5475 (PI: Shara); (8) *HST* GO-9891 (PI: Gilmore).

*The adopted n_k applied to all other observations (see Section 2.4.2).

science targets within cluster tidal radii. For the selection of science targets beyond the cluster tidal radii, the MCPS catalogue, however, was exclusively used. For NGC 1466, NGC 1841, and NGC 2257 (which are not covered by MCPS), we used Jeon et al. (2014)’s catalogue obtained by the Cerro Tololo Inter-American Observatory (CTIO) 0.9 m SMARTS telescope. Once the *Gaia* DR2 catalogue was available, we exclusively selected candidates based on the *Gaia* CMDs and requiring that the parallax of any candidate was ≤ 0.15 mas. The catalogues used for each cluster’s target selection are listed in column 7 of Table 3.

We adopted the positions of targets from different catalogues in the following way. The coordinates for targets selected and observed before *Gaia* DR2 were tied to the NOMAD astrometric system (Zacharias et al. 2004). We cross-matched all stars brighter than 17.5 mag in *V*-band from *HST* images or MCPS catalogue with stars in the NOMAD catalogue, and transformed the coordinates on to the NOMAD frame. For some *HST*-selected stars, rather large astrometric corrections of up to 1–2 arcsec were necessary. For reference, the M2FS fibre apertures are 1.2 arcsec in diameter, and systematic precision of 0.25 arcsec is typically required. After *Gaia* DR2 and to ensure the accuracy of positional dependent kinematic analysis, we cross-matched all targets with *Gaia* DR2 and adopted their J2000 positions.

Fibre positions to sample the backgrounds in and around the clusters were also identified for many, but not all, of the cluster fields

of our sample. Before the release of *Gaia* DR2, background/sky locations were selected by eye within the tidal radii of clusters were identified using F555W *HST* images when available. For regions outside the clusters’ tidal radii, background/sky positions were randomly chosen from the DSS red-band images out to at least 2 arcmin from the clusters’ centres. After the release of *Gaia* DR2, background/sky positions within the clusters were selected by identifying random positions that had no stars within a radius of 3 arcsec in *Gaia* DR2.

2.3 Observations and data reduction

The spectral data used in this study were obtained with the M2FS (Mateo et al. 2012) on the *Magellan*/Clay Telescope over 26 nights during a campaign lasting from 2016 September to 2019 March. The detailed spectral configuration parameters are the same as described in S19. In summary, the single-order spectra ranged from 5130 to 5192 Å in wavelength with a mean resolution of 18 000. In parallel to these single-order observations, we usually also obtained data using a broad order-isolating filter that allowed us to obtain spectra covering 23 orders from 4058 to 5524 Å for up to five targets (typically four stars and one background position). For this study, we use only the same order employed in the single-order spectra from these multi-order data. Table 3 lists the full set of observations, including the observing date, the on-target exposures, the offset exposures, and

the number of background/sky positions (N_{sky}) assigned together with scientific targets. Though not detailed in the table, additional calibration data (e.g. flats, aperture reference spectra, ThArNe arc spectra, twilights, darks, and biases) were obtained throughout the relevant M2FS runs. These calibration data were required by the standard data reduction steps described in S19.

For most clusters we also obtained exposures while the telescope was deliberately offset from the nominal target positions in order to provide another way to determine background contamination. These offset exposures were taken in order to sample the local background for every target throughout the cluster and in the corresponding field. Such offset exposures are not available for all cluster fields (see Table 3). The total exposure times on these offset positions ranged from 10 per cent to 50 per cent of the on-target exposure times (see Table 3 for the actual on- and off-target exposure times for all clusters). We obtained offset background measurements of this sort for 18 of the 29 visits for the 26 clusters in our sample (we visited three clusters – NGC 1978, NGC 2203, and NGC 2209 – on two separate occasions each; see Table 3).

All data were processed using an M2FS pipeline based on IRAF.² The principal end products of this pipeline are the sky-subtracted spectra and their associated variances. In the pipeline, the reduction processes were largely the same as those described in S19, except for the last step – an improved background/sky subtraction – as described in the following subsection.

2.4 Background correction

Background/sky contamination is significant in our data set and challenging to measure, in part due to the presence of unresolved light from the clusters and also due to telluric components arising from sunlight scattered within the atmosphere and reflected off the moon. Together, these sources cause the backgrounds to vary in intensity and spectral character as a function of location relative to cluster centres. Background uncertainties generally have minor impact on the quality of kinematic measurements in the Mgb spectral region used in this study (see Section 2.3). However, good background measurements are required to obtain reliable stellar parameters such as surface gravity and metallicity. Since both of these parameters are used to help determine cluster membership (see Section 4), the precision of background subtraction has indirect impact on the derived dynamical properties of many of the clusters in our sample. In this section, we describe how we sampled the backgrounds in our cluster fields and the methods we developed and tested to apply background corrections to our spectral data.

2.4.1 Background sampling

For most of the clusters in our sample, we assigned from 8 to 16 dedicated sky fibres within the tidal limits of the clusters (see Section 2.2 regarding how background locations were identified). There were some exceptions. NGC 1466 had no pre-determined background positions (see Section 2.2); in this case we plugged 42 unassigned fibres to random open holes in the cluster’s plug plate over as much of the full radial extent of the cluster as possible. For NGC 339 and NGC 419, only 3 and 5 sky fibres, respectively, were assigned within their tidal radii. In the case of NGC 1978, our first

visit (of two) had only one sky-fibre assigned, while the second visit had none (see Section 2.4.2 for more on this system). In all cases, the fairly limited number of sky fibres reflects the relatively large (14 arcsec) minimum spacing between fibres which made it difficult to pack a large number of background fibres in the central regions of the clusters.

In order to sample the background in more locations, we began to obtain observations at offset positions (see Section 2.3). The advantages of this approach are that the local background can be sampled close to each target star (typically about 5 arcsec), and we obtain denser background sampling near the clusters’ centres. The disadvantages are that the offset positions were often contaminated by relatively bright stars, the off-target exposure times were typically only 10–50 per cent the total time on the science targets (see Section 2.3 and Table 3), and the offset exposures were not clearly contemporaneous with target observations which can compromise their utility due to changes in observing conditions (e.g. moonrise/set, airmass changes, onset of twilight).

2.4.2 Background subtraction

In S19, we developed two techniques, dubbed ‘Method A’ and ‘Method B’, to estimate background contributions to M2FS data in and around MC star clusters. ‘Method A’ is suitable for clusters that have dedicated sky fibres and offset-sky observations. Method B can be used when offset-sky observations are not available, only sky fibres which may or (usually) may not sample the inner parts of a cluster adequately. This method requires having a set of Method A clusters available in the sample. Details regarding how these methods were applied in this study can be found in Appendix B.

For the sake of consistency, the final adopted background profiles for our sample were obtained by applying Method B to all clusters regardless of whether a given cluster had both dedicated sky fibres and offset-sky data. We then adopted the ‘median minus $1-\sigma$ ’ profile as the best estimate for a clusters background-light profile. In practice, this profile corresponds to that of NGC 361, the cluster whose background profile is at the 16 percentile (the fifth of 17 profiles; that is, ‘median minus $1-\sigma$ ’) of the rank-ordered distribution of the background profiles.³ The background spectrum for every target star within a given cluster was calculated as the normalized median background spectrum from sky/offset fibres located beyond the tidal radius scaled by the adopted background profile for that cluster.

One cluster in our sample – NGC 1978 – has inadequate data for reliable background subtraction. Specifically, our first visit deployed only one sky-fibre position, and no offset observations were obtained. Our second observation of this cluster lacked both a sky-fibre position and offset observations. As a result, neither Method A nor B be applied in this case. This limits our ability to obtain useful surface gravities and metallicities for the stars in NGC 1978. The kinematic data, however, remain useful, though as we shall see, somewhat enigmatic compared to any other cluster in our sample. We will return to this special case in Section 5.2.

³The use of the ‘median minus $1-\sigma$ ’ background profile was found to produce consistently flat metallicity profiles within the clusters (see Section 4.2 below). In contrast, the median background profiles produce strong inward-rising radial metallicity gradients in nearly all the clusters of our sample. This is precisely what one would expect if the background is being over-subtracted in the central cluster regions. We take this as corroborating evidence that the ‘median minus $1-\sigma$ ’ backgrounds are to be preferred.

²IRAF is a collection of astronomical data reduction software originally written at the National Optical Astronomy Observatory (NOAO).

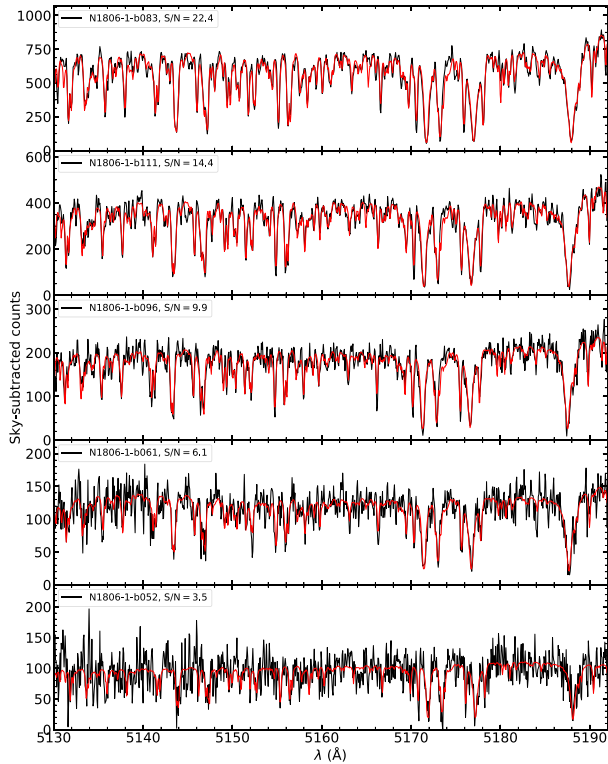


Figure 2. Representative spectra (black) for five stars observed with M2FS in NGC 1806 which conveniently span almost the full range of S/N of usable spectra within our full data set. The spectra shown here have been corrected for backgrounds as described in Section 2.4.2. The red lines are the best-fitting spectral models determined with the Bayesian spectral fitting analysis described in Section 3.1. The legends within each panel list the target ID (following the naming system of Table 5 and the median S/N per pixel of each spectrum).

2.5 Total working sample

Based on the reduction steps described in Sections 2.3 and 2.4, our full M2FS data set consists of 3137 background-subtracted target spectra of 2901 distinct targets in the fields of the 26 MC star clusters we observed for this study (see Table 1). The final total includes the effect of two clusters, NGC 2203 and NGC 2209, where we have combined background-subtracted spectra results, suitably weighted, from independent visits to each cluster (see Table 3). Some representative examples of background-subtracted spectra that span nearly the full range of S/N in our sample are shown in Fig. 2. A full listing of the results from the 2901 spectra is provided in Table 5.

We have expanded this data set with previously published-kinematic data of comparable quality for targets in and near the clusters of our sample (Table 1). We restricted our sources to those with typical velocity precisions of less than about 3 km s^{-1} (Fischer et al. 1992b, 1993; Ferraro et al. 2006; Mucciarelli et al. 2008; Mucciarelli et al. 2010, 2014; Mackey et al. 2013; Patrick et al. 2020). Table 4 provides some details of the data obtained from these sources, including the total number of targets previously studied and the number of stars in common with our M2FS sample. The table also lists the cluster systemic velocities from the previous studies and the difference relative to the corresponding systemic velocities obtained from the new M2FS data in the sense M2FS results minus previously published results. Details regarding how we combined earlier and M2FS data are described in Section 4.3, and final consolidated results are incorporated in Table 5. We critically compare repeat

measurements of common stars in Section 5.3. With the addition of these earlier data sets, our final comprehensive sample consists of 3376 spectroscopic results for 3095 distinct sources in our target clusters.

3 SPECTRAL ANALYSIS

3.1 Bayesian fitting of M2FS spectra

Our analysis of the background-corrected spectra employed the same Bayesian formalism described in Walker, Olszewski & Mateo (2015a), Song et al. (2017) and adopted in S19. The outputs of this analysis include estimates of line-of-sight (LOS) velocity, surface gravity, and metallicity for every star along with associated uncertainties. Some examples of the best-fitting spectra obtained with this method are shown in Fig. 2.

For this study, we altered one important aspect of the analysis compared to S19. Namely, the effective temperatures of all stars were treated as priors based on the $G_{\text{BP}} - G_{\text{RP}}$ colours provided in the *Gaia* DR2 rather than from $V - I$ colours from, for example, *HST* images (the method we used for NGC 419 and NGC 1846 in S19). The resulting temperatures were then forced to remain fixed throughout the spectral-fitting process. The details of this procedure are described in the following section.

3.2 Effective temperature priors

As argued in Song et al. (2017), an effective temperature prior (T_{eff}) helps us to break the temperature–metallicity degeneracy in the posterior probability density functions (PDFs) calculated using our Bayesian analysis (see Section 3.1). This degeneracy affects our data strongly due to the limited wavelength range (5130 to 5192 Å) of the single-order spectra we obtained (see Section 2.3). To help mitigate this problem, our application of the Bayesian spectral fitting procedure fixes the effective temperature throughout the optimization process for a given star.

One complication with this approach is that if we were to use the *Gaia* DR2 colours to estimate T_{eff} directly, our temperature estimates would likely exhibit significant systematic and random uncertainties. There are two reasons for this. First, many of the stars in our sample are comparatively faint ($G > 17 \text{ mag}$) for the *Gaia* sample, and so their formal photometric errors are moderate (Evans et al. 2018). Secondly, given the high source densities in many of our fields, background variations and contamination due to crowding can lead to significant additional uncertainty in the *Gaia* $G_{\text{BP}} - G_{\text{RP}}$ colours (Maíz Apellániz & Weiler 2018; Weiler 2018). Moreover, in clusters, blending of red giants with typically hotter, fainter cluster members tends to drive colours systematically to the blue.

These issues are particularly problematic for T_{eff} determinations in RGB stars since even a small shift in colour can lead to a significant temperature change which, in turn, degrades the precision of the metallicity estimate of a star. To mitigate these problems, we have devised a method to estimate T_{eff} that depends primarily on the observed G magnitude of a star and that relies on the fact that the stars in our sample were almost exclusively selected to be located on the RGBs of their respective clusters (see Section 2.2). This methodology could not be applied to NGC 330, NGC 458, and NGC 1850 as they lack an extended RGB owing to their young ages.

The method starts by identifying an isochrone that matches the adopted age and metallicity of a given cluster (see Table 1). For this study we have chosen to use the PARSEC isochrones (version 1.2S; Bressan et al. 2012) converted to the *Gaia* DR2 photometric system

Table 4. Previous kinematic data of clusters in our sample^a.

Galaxy	Cluster	N_{sample}	N_{common}	$V_{\text{sys,prev}}$ (km s^{-1})	ΔV_{sys} (km s^{-1})	Source	Code
(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)
SMC	NGC 330	16	5	153.7 ± 1.0	-0.87 ± 1.27	Patrick et al. (2020)	Pa20
LMC	NGC 1783	6	0	277.6 ± 1.0	1.98 ± 1.02	Mucciarelli et al. (2008)	Mu08
LMC	NGC 1806	8	2	228.6 ± 0.5	0.93 ± 0.87	Mucciarelli et al. (2014)	Mu14
LMC	NGC 1846	105	17	239.1 ± 1.0	0.21 ± 0.46	Mackey et al. (2013)	Ma13
LMC	NGC 1850	52	2	251.4 ± 2.0	-2.43 ± 2.16	Fischer et al. (1993)	Fi93
LMC	NGC 1978	35	8	293.3 ± 1.0	-0.19 ± 1.11	Fischer et al. (1992b)	Fi92
LMC	NGC 1978	11	4	293.1 ± 0.8	0.01 ± 1.03	Ferraro et al. (2006)	Fe06
LMC	NGC 2257	6	1	299.4 ± 1.5	2.42 ± 1.53	Mucciarelli, Origlia & Ferraro (2010)	Mu10

Note. ^aColumns 3 and 4 list the number of stars in the studies listed in column 7 with data of sufficient kinematic precision for inclusion in our current sample (N_{sample}) and the number of stars in common with our current sample (N_{common}), respectively. Columns 5 and 6 list the systemic cluster velocities for each cluster from the cited sources and the difference in the sense $\Delta V_{\text{sys}} = V_{\text{sys,M2FS}} - V_{\text{sys,prev}}$. Column 8 lists for each source cited in Column 7 a short code used in Table 5.

following the prescription of Maíz Apellániz & Weiler (2018).⁴ We then plotted the synthetic isochrones on to the observational G versus $G_{\text{BP}} - G_{\text{RP}}$ plane (an example is shown in the left-hand panel of Fig. 3). In all cases, we shifted the isochrones to the observational plane with the appropriate reddening values and distance moduli for each cluster (Table 1). As suggested by the PARSEC model website, these corrections used the extinction parameters $A_G = 0.86A_V$, $E(G_{\text{BP}} - G_{\text{RP}}) = 0.42A_V$ and $R_V = 3.1$ within the extinction relations from Cardelli, Clayton & Mathis (1989) and O’Donnell (1994). For computational convenience, we then fit each isochrone in the observational plane to a third-order spline to smooth out round-off noise and some of the non-monotonic evolutionary behaviour in the isochrones themselves. An example of such a fit is illustrated in the left-hand panel of Fig. 3.

The *Gaia* DR2 photometry for every cluster was corrected to the ‘true’ G -band system as suggested by Anders et al. (2019) before plotting the data in the observational CMD with their respective isochrones. Each star in our spectroscopic sample was then projected at constant G magnitude on to the spline fits to the RGBs on a cluster-by-cluster basis. The corresponding $G_{\text{BP}} - G_{\text{RP}}$ colour at that point in the isochrone was adopted as ‘the’ colour for that star. The middle panel of Fig. 3 provides a schematic illustration of this process, and was adopted for all clusters older than about 0.5 Gyr since their RGB sequences could be reasonably defined from the models.

For the three clusters in our sample younger than 0.5 Gyr (NGC 330, NGC 458 and NGC 1850), we did not correct the colours in this manner but simply adopted the published *Gaia* DR2 colours corrected as prescribed by Anders et al. (2019). For these systems, the red (super)giants tend to be brighter than 17 mag in G -band and hence exhibit relatively small colour errors, and these stars are less susceptible to crowding/background-related photometric errors. Moreover, these younger stars tend to cluster in a region in the CMD rather than on a distinct giant branch (see, e.g. Mermilliod 1981; Alcaino et al. 2003; Correnti et al. 2017; Milone et al. 2018); hence they are ill-suited for the temperature-determination procedure we used for the rest of our sample which assumes a well-defined red giant locus with little colour spread at a given luminosity.

Fig. 4 shows a histogram of the colour shifts in $\Delta(G_{\text{BP}} - G_{\text{RP}})$ that we determined for every star with a tabulated *Gaia* DR2 colour

using the method described above. The mode of this distribution is offset from, but close to zero, confirming that most of these objects are consistent with being RGB/AGB stars. This supports the efficacy of our approach. Since the slope of the giant branch is approximately $\delta(G_{\text{BP}} - G_{\text{RP}})/\delta G \sim 0.2$, the colour error is about 20 per cent the G -magnitude error for a given star if it is on the RGB as assumed. Consequently, the resulting error distribution of the adopted colours is considerably narrower than the distribution of the ‘raw’ $\Delta(G_{\text{BP}} - G_{\text{RP}})$ implied by Fig. 4. It is also evident in Fig. 4 that the $\Delta(G_{\text{BP}} - G_{\text{RP}})$ distribution exhibits a moderate tail to positive values. This implies that some comparatively hot stars fail our assumption that they are on the RGB and hence they have been assigned colours that are systematically too red.

The fact that the mean offset in $\Delta(G_{\text{BP}} - G_{\text{RP}})$ differs between the LMC and SMC (Fig. 4) suggests that our colour corrections may depend on metallicity (larger effect for larger metallicities in the LMC). We will return to this problem in Sections 4.2 and 5.4 where we use Fig. 4 to help us determine mean cluster metallicities.

The final step in estimating T_{eff} values for our target stars was to convert their $G_{\text{BP}} - G_{\text{RP}}$ colours – either the ‘raw’ values for stars in the three young clusters or the model-based values determined using the procedure above – to temperatures. We again used the PARSEC RGB sequence for each cluster (appropriately reddened and adopting the ages and metallicities in Table 1) to define a colour–temperature relation for each case (see the right-hand panel of Fig. 3 for an example). The final T_{eff} values were then tabulated for every star and these values are provided in Table 5.

3.3 Velocity uncertainty correction

To reliably measure small internal velocity dispersion of systems such as the star clusters in our sample, high-quality kinematic data and reliable error estimates are essential (see e.g. Kamann et al. 2016, S19). In S19, we analysed repeat measurements of individual stars to empirically assess the quality of the uncertainties computed by the Bayesian spectral analysis for the cases of NGC 419 and NGC 1846. We found that individual velocity uncertainties returned by the Bayesian analysis underestimated the true uncertainties by approximately 23 per cent and 12 per cent, respectively.

In this study, we explore this question again for the two clusters in our sample with the lowest derived central velocity dispersions, NGC 2155 and SL 663 (see Section 4). As in S19, we compared the velocities from individual exposures, s_i , and then fit a Gaussian to the error distribution expressed in units of the formal error, σ_i , returned

⁴In practice, we computed tailor-made isochrones for each cluster in the appropriate photometric system using the PARSEC website (<http://stev.oap.d.inaf.it/cgi-bin/cmd>).

Table 5. Sample of all stars. (The full table is available online as supplementary material.)

Galaxy	Cluster	ID	RAJ2000 (deg)	DEJ2000 (deg)	G (mag)	$G_{BP} - G_{RP}$ (mag)	T_{eff} (K)	$Gaia$ DR2 ID	S/N ^a	v_{los} (km s ⁻¹)	log g (dex)	[Fe/H] _{raw} (dex)	P_M	P'_M	Flag ^b	Source ^c
(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)	(9)	(10)	(11)	(12)	(13)	(14)	(15)	(16)	(17)
SMC	Kron 3	K3-1-b002	6.420511	-72.862261	18.35	1.20	4869	4688782811694646272	4.1	111.20 ± 0.78	1.63 ± 0.21	-1.03 ± 0.08	0.00	0.00	0000000	M2FS
SMC	Kron 3	K3-1-b004	6.366614	-72.850370	18.29	1.26	4854	4688782983493342336	4.3	130.36 ± 0.49	1.35 ± 0.20	-0.88 ± 0.08	0.00	0.00	0000000	M2FS
SMC	Kron 3	K3-1-b005	6.307460	-72.842177	17.28	1.45	4577	4688783022152125440	7.3	160.47 ± 0.31	1.21 ± 0.09	-0.75 ± 0.04	0.00	0.00	0000000	M2FS
LMC	NGC 1978	N1978-3-r049	82.190452	-66.235682	16.14	1.04	3840	4660340443101129856	4.4	292.97 ± 0.52	0.51 ± 0.10	-0.89 ± 0.07	1.00	1.00	0000010	M2FS
LMC	NGC 1978	N1978-3-b034	82.209310	-66.256802	15.94	1.90	3755	4660340271299657216	13.5	292.38 ± 0.16	0.84 ± 0.04	-0.65 ± 0.04	0.99	0.99	0000000	M2FS+Fi92+Mu08
LMC	NGC 1978	N1978-3-b037	82.211433	-66.245587	15.76	1.95	3755	4660340370035622144	14.7	295.43 ± 0.26	0.84 ± 0.05	-0.66 ± 0.04	1.00	1.00	0000000	M2FS+Fi92
LMC	NGC 1978	06	82.193775	-66.236033	-	-	-	-	-	291.21 ± 1.30	-	-	1.00	1.00	0000000	Fi92
LMC	NGC 1978	07	82.192219	-66.238693	-	-	-	-	-	288.76 ± 0.20	-	-	1.00	1.00	0000000	Fi92+Mu08

^aNotes. ^aMedian S/N per pixel of M2FS spectrum.

^bThese seven-digit flags denote with a '1' the following: Rejection due to poor skew/kurtosis values in the Bayesian spectra fits (Digit 1); Excessive velocity error (Digit 2); Carbon star (Digit 3); Foreground dwarf (Digit 4); Member of a non-cluster/non-MC population (Digit 5); Has a large colour offset for T_{eff} determination (Digit 6); Likely metallicity non-member (Digit 7). Details about how these flags are set can be found in Section 3.4.1, Section 3.4.2, Section 3.4.3, Section 4.1, and Section 4.2.

^cSources for the LOS velocities. M2FS denotes this work. Other codes are listed in column 4 of Table 4.

by the Bayesian analysis for the i -th spectrum (see Kamann et al. 2016 for analogous analysis). One complication with this test is that there are small systematic velocity shifts in M2FS data that correlates with the temperature of the instrument near the fibre pseudo-slit (see Walker et al. 2015b). The spectra from the individual exposures for both clusters were first corrected for these shifts and referenced to a common velocity scale before generating the s/σ distribution. The results of this analysis show that the velocity uncertainties for these clusters were similarly underestimated as found in S19, i.e. approximately 16 per cent and 18 per cent for NGC 2155 and SL 663, respectively (see Section 3.1.1 and fig. 4 of S19 for details).

Together with our earlier results from S19, we conclude that our application of the Bayesian spectral analysis underestimates the true velocity errors by 17 per cent ± 3 per cent. Consequently, we have increased the velocity uncertainties from all Bayesian spectral analysis results by 17 per cent prior to carrying out any dynamical analyses of these systems (see Sections 4 and 5).⁵

3.4 Towards a final M2FS spectral sample

3.4.1 Rejecting spectra with poorly determined parameters

Our Bayesian analysis as described in Section 3.1 generates posterior parameter distribution functions (PDFs) for 14 free parameters (the 15th, T_{eff} , is fixed as described in Section 3.2) used to characterize the model spectra (see Song et al. 2017 for details). The analysis measures the skew, S , and kurtosis, K , of all of these PDFs. We use the S - K values for the PDF of the LOS velocity (v_{los}) parameter for every spectrum to carry out an initial quality cut on the data. After inspection of the distribution of the skew/kurtosis values for the LOS velocities, we found that the same cuts in S and K from Walker et al. (2015a) were suitable for our data, namely $|S| > 1$, or $|K - 3| > 1$. This 'SK cut' led to the rejection of 71 spectra from our sample. These cases are flagged in column 17 of Table 5; see Appendix C1 for further details. We will not use these stars in any subsequent analyses in this paper.

3.4.2 Stars with anomalous velocity uncertainties

Fig. 5 shows the distribution of corrected velocity uncertainties, ε_v , as a function of median S/N ratio for every pixel in every target spectrum obtained for this study in log-log space. The well-populated ridge lines reveal a robust global trend of velocity error with S/N for our M2FS data set. The approximately linear trend visible in the log-log space indicates that a good fit to these data has the form of a power law.

With the aim of using the relation between ε_v and S/N to remove outliers from our kinematic sample, we fit a linear relation in the log-log plane using an unweighted least-squares fit to all points corresponding to spectra with median S/N ≤ 20. Targets removed from the SK-cut described in Section 3.4.1 were not used to compute

⁵The origin of the 17 per cent correction likely lies in the fact that our data-reduction pipeline rebins the data during the extraction and wavelength-calibration steps. This is confirmed by independent reductions of similar data by M. G. Walker (private communication) who finds that unbinned analyses lead to a much smaller correction consistent with unity. His results also independently confirm the velocity-uncertainty correction factor of 17 per cent that we obtained. We consider the correction to be sufficiently small and sufficiently well-determined to use it to adjust our formal velocity error estimates for the purposes of this study.

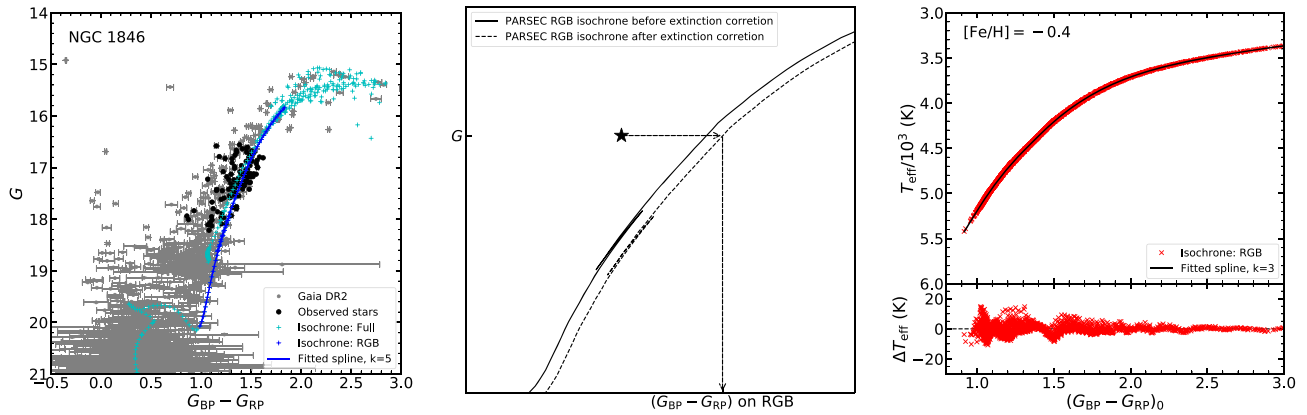


Figure 3. Left-hand panel: The colour-magnitude diagram for NGC 1846 based on *Gaia* DR2 and using the DR2 photometric system. A PARSEC model isochrone (Bressan et al. 2012) using cluster’s age and metallicity estimates (see Table 1) is overplotted (light blue crosses). Gray dots correspond to all *Gaia* DR2 stars located 2 arcmin from the cluster centre, while black dots represent the stars observed in our sample. The blue crosses highlight the RGB sequence of the isochrone, and the blue solid curve is a spline fit to this RGB segment. (Middle panel) Following the discussion in Section 3.2, this plot illustrates schematically how a given star in the CMD is projected to the adopted RGB sequence to determine the star’s model-based $G_{BP} - G_{RP}$ colour. Right-hand panel: The red symbols show an example of the *Gaia* DR2 colour/temperature relations obtained from the PARSEC models. The relation shown here is suitable for the clusters, such as NGC 1846, with a mean metallicity of $[Fe/H] \sim -0.4$ (see Table 1). These individual colour-temperature points were fit to a third-order spline (black line; fit residuals are shown in the lower subpanel). We used this spline fit to convert the adjusted colours (see middle panel) to effective temperature for each star. These temperatures were fixed during the Bayesian spectral fitting procedure (see Section 3.1).

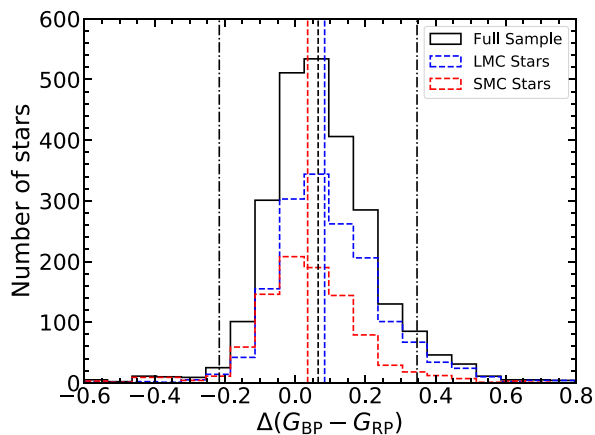


Figure 4. A histogram of the colour shifts, $\Delta(G_{BP} - G_{RP})$, between the *Gaia* DR2 colour of a given star in our data set and our adjusted colour obtained as described in Section 3.2 and Fig. 3. For 2543 stars in the full sample, the biweight mean offset $C_{BI,dcolour}$ in $\Delta(G_{BP} - G_{RP})$ (denoted by the dashed line) is 0.065 with a $2S_{BI,dcolour}$ range of ± 0.281 (denoted by the dot-dash lines). A total of 203 stars are located outside the dashed lines, though the distribution is clearly skewed towards positive values (68 stars on the negative side and 135 on the positive side). This likely represents the fact that some stars located bluewards of the adopted RGB sequences in individual clusters (see the middle panel of Fig. 3) are hotter than red giants. For the purpose of metallicity determination described in Section 4.2, we use the stars within two biweight dispersion scales of the biweight mean, i.e. $|\Delta(G_{BP} - G_{RP}) - C_{BI,dcolour}| \leq 2S_{BI,dcolour}$. This ensures that only stars with relatively small colour shifts – likely true RGB stars – are used to determine the cluster metallicity. The blue and red dashed histograms show the distributions in $\Delta(G_{BP} - G_{RP})$ for the LMC (1595 stars) and the SMC (948 stars). The mean offsets of these histograms are $C_{BI,dcolour} \pm 2S_{BI,dcolour} = 0.084 \pm 0.283$ for the LMC and 0.036 ± 0.269 for the SMC.

this fit. The fitted relation is shown as a solid straight line in the log–log plot (the coefficients for the fit equations are provided in Fig. 5). We then computed logarithmic residuals ($\Delta = \log_{10} \varepsilon_v - \log_{10} \varepsilon_{v,fit}$) to estimate the Tukey’s biweight location ($C_{BI,\Delta}$) and biweight scale

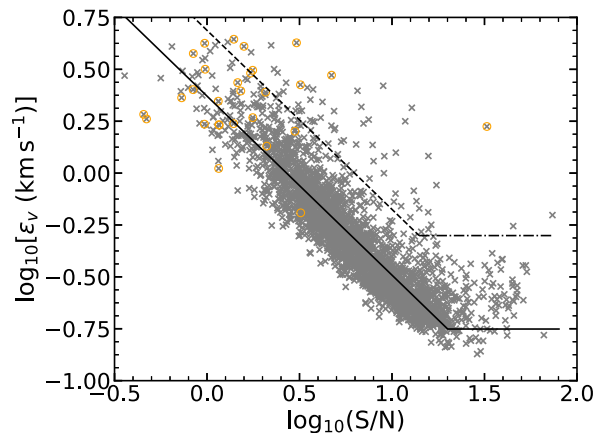


Figure 5. Corrected LOS velocity uncertainties versus median S/N per pixel for all stars from the 26 clusters in log–log plane. The solid lines correspond to a linear relation in the log–log plane fit to all points with $S/N \leq 20$ (excluding the yellow circles rejected by the SK discussion in Section 3.4.1). For $S/N > 20$, we assume ε_v is constant at 0.18 km s^{-1} . The dashed and dot-dashed lines together denote the upper acceptance limits in these plots as described in Section 3.4.2.

($S_{BI,\Delta}$) of the data about the fitted line in the log–log plot (Beers, Flynn & Gebhardt 1990). The dashed straight line in the log–log plot in Fig. 5 indicates the line offset upward by $C_{BI} + 3S_{BI}$ from the fitted relation; the dashed line in the linear plot (left-hand panel) shows the power-law relation corresponding to this offset line. We also chose to keep all stars with high-S/N spectra that resulted in $\varepsilon_v \leq 0.5 \text{ km s}^{-1}$; this condition is shown as a dot-dash line in both panels.

We define the dashed/dot-dashed lines in Fig. 5 as a ‘rejection boundary.’ Excluding the 71 stars already rejected as SK outliers (Section 3.4.1), 62 additional stars lie above this boundary. We inspected each of these 62 spectra and found that we could classify them into distinct categories, each of which are flagged in column 17

of Table 5. Details of the classifications of these spectra are provided in Appendix C2.

In S19 we identified eight stars in NGC 419 and NGC 1846 (four in each) that we rejected using similar criteria as described here but based only on the small samples of these two clusters. The four stars in our NGC 419 sample (three C stars and one blended source) rejected in that paper have also been rejected in the present sample. However, the four stars in our NGC 1846 sample that were rejected in S19 consist of three weak C stars and one mildly blended source. All four of these stars were not rejected in the present analysis using our full spectral sample. This very slightly – and statistically indistinguishably – alters our final dynamical results below (see Section 5) for NGC 1846 compared to the results reported in S19.

3.4.3 Foreground dwarfs

Surface gravity is one of the parameters returned by the Bayesian analysis (Section 3.1) and this can be used to identify foreground dwarf stars in our sample. As seen in Fig. 6, there is a clear tail in the distribution of $\log g$ extending to larger values. We have adopted $\log g \geq 3.2$ as the dividing value between giants (lower $\log g$) and dwarfs. A total of 84 stars in our sample were removed using this criterion and are flagged in column 17 of Table 5. Representative spectra of these stars are illustrated in Appendix C3.

At this stage, our Bayesian spectral-fitting results have allowed us to produce a well-defined M2FS sample consisting of 2680 spectra of (mostly) RGB stars in our 26 target clusters (see Table 5). Fig. 2 presents some representative spectra of normal stars in our sample along with their Bayesian spectral fits.

4 CLUSTER MEMBERSHIP ANALYSIS

We closely follow the expectation–maximization (EM) analysis used in S19 to determine simultaneously the cluster dynamics and the cluster membership probabilities of all stars in our samples. For most of the clusters, we assume in the EM analysis that there are two stellar populations in each kinematic cluster sample: one corresponding to the cluster, and one corresponding to the LMC/SMC field stars. Following S19, we still adopt single-mass K66 models⁶ (as listed in Table 2) to generate the projected velocity dispersion profile of each cluster, while the field population is represented by a Gaussian distribution. In our EM analysis for two old LMC clusters, NGC 1466 and NGC 1841, however, we only assume a single population (also in the form of the K66 model) because both clusters are so distant from their parent galaxy (the LMC) that they lack significant numbers of LMC field stars.

An issue that arises at this stage is that most clusters contain stars that are either associated with a third population not included in our model (see equation 6 of S19), or stars that have metallicities that deviate significantly enough from their respective cluster mean metallicity to call in question their membership (keeping in mind that metallicity is not used directly in assigning membership probabilities in our EM analysis; see S19). We describe here how we have used preliminary EM analysis results and metallicity estimates from the Bayesian spectral fitting (Section 3.1) to develop a final sample for kinematic analysis. We also describe how we have supplemented this final sample with results from previous studies.

⁶The K66 models are realized through LIMEPY (Gieles & Zocchi 2015) as described in S19.

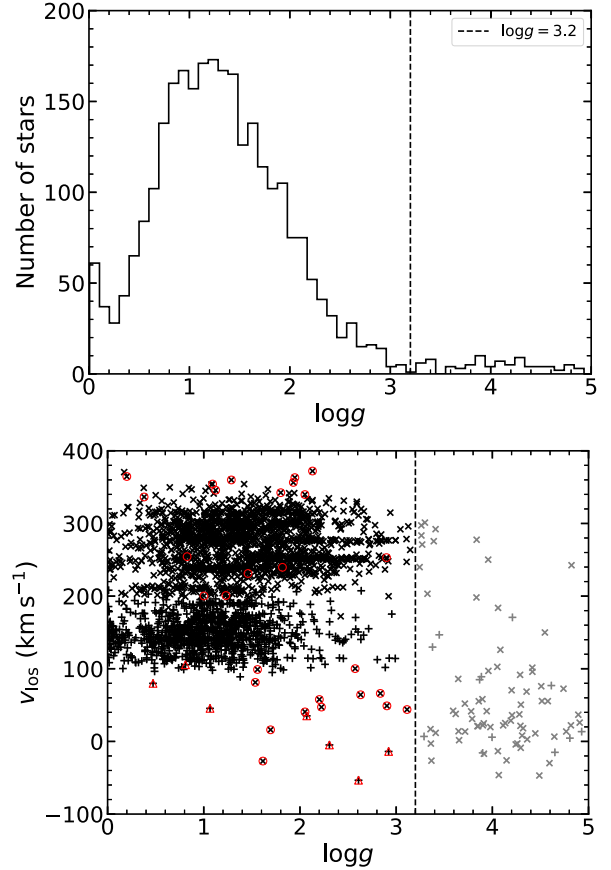


Figure 6. Top panel: A histogram of the surface gravity for 2764 stars in our M2FS sample after removing the stars with either poorly determined parameters (see Section 3.4.1) or anomalous velocity uncertainties (see Section 3.4.2). The vertical dashed line corresponds to an adopted empirical criterion, $\log g = 3.2$, to divide giants and (foreground) dwarfs, as described in Section 3.4.3. The 84 stars to the right of dashed line were removed from our sample as likely foreground dwarfs. Bottom panel: A plot on the $v_{\text{los}}-\log g$ plane of 2764 stars in our M2FS sample after removing the stars with either poorly determined parameters (see Section 3.4.1) or anomalous velocity uncertainties (see Section 3.4.2). Crosses denote stars observed in LMC fields, while plus signs refer to stars observed in SMC fields. The dashed vertical line correspond to the criterion, $\log g = 3.2$, used in Section 3.4.3 to separate a total of 84 (foreground) dwarfs (grey crosses and plus signs) from giants (black crosses and plus signs) in our sample. The red circles (LMC) and red triangles (SMC) denote the 35 rejected stars that we concluded are associated with a third population as discussed in Section 4.1.

4.1 Removing stars from a third stellar population

Stars from a stellar population not explicitly included in our EM model can distort the estimation of the field population and hence the membership probability of other stars in the sample. We flagged such stars on a cluster-by-cluster basis using the following criterion:

$$\frac{\Delta v_{\text{los}}}{\sigma} \equiv \frac{|v_{\text{los},i} - \overline{v_{\text{los}}}|}{\sqrt{\varepsilon_{v_{\text{los},i}}^2 + \sigma_{v_{\text{los}}}^2}} \geq 3, \quad (1)$$

where $v_{\text{los},i}$ and $\varepsilon_{v_{\text{los},i}}$ are the LOS velocity and its associated uncertainty of the i -th star in a cluster sample, respectively, while $\overline{v_{\text{los}}}$ and $\sigma_{v_{\text{los}}}^2$ are weighted average velocity and the corresponding weighted standard deviation of the corresponding cluster sample. Using this process, a total of 35 stars in 17 cluster samples were identified as third-population contaminants (bottom panel in Fig. 6);

all are flagged in column 17 of Table 5. The generally low surface gravities of these rejected stars suggests they are likely Galactic halo giants or giants associated with an extended halo population of the MCs, neither of which were explicitly accounted for in our EM model (see Section 4 and S19).

Using the remaining sample of 2645 stars, we reran the EM analysis for every cluster to assign a preliminary membership probabilities, $P_{M,i}$, to each star.

4.2 Removing stars with anomalous metallicities

We can also use metallicities obtained from the Bayesian analysis (see Section 3.1) to identify stars with chemical abundances that differ significantly from the mean of their respective cluster’s [Fe/H] distributions. Such deviant metallicities may indicate non-membership, but they could also result from issues such as poor background subtraction or poor temperature assignments.

We applied two criteria to identify potential metallicity non-members. First, the preliminary membership probabilities had to satisfy the condition $P_{M,i} \geq 0.5$. Secondly, the colour shifts, $\Delta(G_{BP} - G_{RP})$ to assign effective temperature to each star (see Section 3.2) had to satisfy the condition $|\Delta(G_{BP} - G_{RP}) - C_{BI,dcolour}| \leq 2S_{BI,dcolour}$, where $C_{BI,dcolour}$ and $S_{BI,dcolour}$ are the biweight mean offset and dispersion scale determined in Section 3.2, respectively. This criterion was designed to avoid metallicity offsets due to systematically invalid temperature estimates (see Section 3.2 and Fig. 4 for details. Stars identified in this manner as likely metallicity outliers are flagged in column 17 of Table 5.

Using all stars satisfying these two criteria, we then calculated for each cluster an initial pair of metallicity parameters – the weighted mean metallicity ($\overline{[Fe/H]}_{SC}$) and weighted standard deviation ($\sigma_{[Fe/H],SC}$). Following this process, probable metallicity non-members were identified using the criterion

$$\frac{\Delta[Fe/H]}{\sigma} \equiv \frac{|[Fe/H]_i - \overline{[Fe/H]}_{SC}|}{\sqrt{\varepsilon_{[Fe/H],i}^2 + \sigma_{[Fe/H],SC}^2}} \geq 2, \quad (2)$$

where $[Fe/H]_i$ and $\varepsilon_{[Fe/H],i}$ are the metallicity and its uncertainty of the i -th star in the EM sample, respectively. This procedure resulted in the removal of 51 stars from 22 clusters, and each is flagged in column 17 of Table 5. The number of stars that satisfy equation (2) (i.e. metallicity non-members) are listed for each cluster in column 3 of Table 6, while the total number that do not satisfy this criterion are listed in column 4. The metallicities listed in columns 5–6 of Table 6 are determined using methods describe below in Section 5.4. After applying this metallicity criterion, a total of 2594 stars in 26 clusters remain in our M2FS kinematic sample.

4.3 Combining previous samples

We have identified all previously published individual stellar velocity measurements with a typical single-star precision better than 3 km s^{-1} (see also Table 4) associated with the clusters in our sample. In this section, we identify and compare stars common to our sample and samples from previous studies with the aim of expanding our final kinematic sample.

We first compared published coordinates of individual stars with those of stars associated with our cluster sample. Stars with coordinates agreeing to within a separation of 1 arcsec or less were considered to be the same star. There were no ambiguous cases in which more than one star from previously published catalogue satisfied this positional matching tolerance. For NGC 1978, we also

Table 6. Mean cluster metallicity on the CG97 scale.

Galaxy	Cluster	N_{rem}	$N_{[Fe/H]}$	[Fe/H] (dex)	$\sigma_{[Fe/H]}$ (dex)
(1)	(2)	(3)	(4)	(5)	(6)
SMC	Kron 3	1	41	−0.96	0.15
SMC	Lindsay 1	6	80	−0.98	0.13
SMC	NGC 152	2	22	−0.73	0.11
SMC	NGC 330	1	4	−0.65	0.10
SMC	NGC 339	8	35	−1.01	0.17
SMC	NGC 361	0	20	−0.75	0.17
SMC	NGC 411	2	19	−0.66	0.09
SMC	NGC 416	0	16	−0.80	0.17
SMC	NGC 419	2	35	−0.66	0.15
SMC	NGC 458	0	14	−0.70	0.20
LMC	Hodge 4	2	30	−0.49	0.12
LMC	NGC 1466	1	22	−1.40	0.16
LMC	NGC 1751	1	20	−0.46	0.14
LMC	NGC 1783	2	52	−0.54	0.10
LMC	NGC 1806	1	27	−0.53	0.12
LMC	NGC 1831	3	49	−0.41	0.15
LMC	NGC 1841	3	64	−1.96	0.12
LMC	NGC 1846	2	36	−0.49	0.08
LMC	NGC 1850	0	26	−0.31	0.20
LMC	NGC 1978	3	37	−0.49	0.10
LMC	NGC 2121	3	38	−0.54	0.11
LMC	NGC 2155	3	36	−0.59	0.12
LMC	NGC 2203	2	64	−0.45	0.12
LMC	NGC 2209	1	50	−0.52	0.15
LMC	NGC 2257	1	57	−1.64	0.11
LMC	SL 663	1	19	−0.51	0.11

cross-matched the Fischer et al. (1992b) and Ferraro et al. (2006) samples to identify any stars in common. These two samples were found to have nine stars in common, three of which are also found in our M2FS sample.

Before combining results from different samples, we measured the systemic velocity offsets for clusters in common among the various data sets. These offsets were computed by subtracting the published systemic velocities (see column 6 in Table 4) from our preliminary EM results (see Section 4.1), and the uncertainties in the offsets were estimated by using the sum of errors in quadrature. When combining results from different samples, we then applied the offsets to the previously published velocities while keeping the published velocity uncertainties unchanged. In all cases, we adopted the M2FS velocity zero points for the velocity scale (Table 7). For each star with multiple measurements, the weighted mean velocities and their weighted errors were calculated after applying these offsets. This process allowed us to increase the sample sizes for seven clusters (see Table 4).

Most of the stars matched by position as described above agree well in velocity. However, two stars have velocity differences three times greater than their combined total velocity errors. The star in NGC 1846, whose detail had been discussed in Section 3.1.2 in S19, is a likely binary. The other star, in NGC 330, is a well-known binary system identified in earlier studies (see e.g. Patrick et al. 2020, and the references therein). Both stars were removed from our final kinematic sample. For the NGC 1846 binary, there are too few measurements to determine a reliable systemic velocity, while for the NGC 330 binary, its mean velocity and metallicity, though both somewhat poorly determined, make it a likely non-member of the cluster. Indeed, for this reason, this star had already been flagged for removal in Section 4.2.

Table 7. Results of the EM analysis.

Galaxy	Cluster	N_{total}	N_{tidal}	V_{sys} (km s^{-1})	$\sigma_{p,0}$ (km s^{-1})	V_{field} (km s^{-1})	σ_{field} (km s^{-1})
(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)
SMC	Kron 3	116	84	132.7 ^{+0.3} _{-0.4}	2.1 ^{+0.2} _{-0.3}	143.3 ^{+3.2} _{-2.4}	23.1 ^{+1.3} _{-1.5}
SMC	Lindsay 1	99	92	140.5 ^{+0.2} _{-0.2}	1.8 ^{+0.2} _{-0.3}	137.7 ^{+3.3} _{-4.4}	16.1 ^{+1.7} _{-2.3}
SMC	NGC 152	107	86	172.4 ^{+0.5} _{-0.9}	2.8 ^{+1.2} _{-1.8}	150.7 ^{+4.0} _{-2.5}	25.2 ^{+1.3} _{-2.4}
SMC	NGC 330	47	28	153.0 ^{+0.7} _{-0.7}	2.4 ^{+0.3} _{-0.7}	138.5 ^{+3.0} _{-3.3}	16.0 ^{+1.7} _{-2.6}
SMC	NGC 339	94	65	112.9 ^{+0.4} _{-0.3}	1.8 ^{+0.2} _{-0.3}	161.1 ^{+2.2} _{-2.5}	18.0 ^{+1.6} _{-2.3}
SMC	NGC 361	119	70	170.3 ^{+0.9} _{-0.9}	4.0 ^{+0.7} _{-1.4}	138.6 ^{+1.3} _{-2.8}	20.1 ^{+1.0} _{-1.2}
SMC	NGC 411	117	56	163.8 ^{+4.5} _{-0.3}	1.7 ^{+16.8} _{-0.6}	147.7 ^{+1.4} _{-3.5}	23.3 ^{+0.8} _{-2.2}
SMC	NGC 416	114	40	155.0 ^{+1.0} _{-0.5}	3.4 ^{+0.6} _{-0.7}	149.2 ^{+2.3} _{-2.9}	27.2 ^{+1.5} _{-1.9}
SMC	NGC 419	110	81	189.9 ^{+0.3} _{-0.2}	2.2 ^{+0.3} _{-0.3}	159.8 ^{+2.7} _{-2.6}	21.8 ^{+1.4} _{-1.7}
SMC	NGC 458	79	24	149.0 ^{+0.8} _{-0.9}	3.5 ^{+0.7} _{-2.2}	146.2 ^{+1.8} _{-1.1}	13.3 ^{+1.5} _{-1.3}
LMC	Hodge 4	112	112	312.7 ^{+0.6} _{-1.3}	2.1 ^{+5.1} _{-0.6}	299.8 ^{+2.0} _{-5.6}	21.8 ^{+3.5} _{-3.9}
LMC	NGC 1466	27	25	202.5 ^{+0.5} _{-0.5}	3.6 ^{+0.4} _{-0.6}
LMC	NGC 1751	113	113	240.4 ^{+0.7} _{-0.7}	2.5 ^{+0.6} _{-0.7}	247.1 ^{+3.4} _{-2.7}	27.4 ^{+3.0} _{-3.1}
LMC	NGC 1783	111	111	279.6 ^{+0.2} _{-0.2}	2.0 ^{+0.2} _{-0.2}	281.9 ^{+3.9} _{-3.5}	24.5 ^{+2.3} _{-2.4}
LMC	NGC 1806	120	120	229.6 ^{+0.4} _{-0.4}	2.4 ^{+0.4} _{-0.6}	265.7 ^{+3.9} _{-3.0}	27.0 ^{+1.7} _{-2.3}
LMC	NGC 1831	102	95	276.8 ^{+0.2} _{-0.2}	1.9 ^{+0.3} _{-0.3}	284.8 ^{+2.8} _{-2.5}	17.5 ^{+1.9} _{-3.0}
LMC	NGC 1841	69	64	210.8 ^{+0.3} _{-0.3}	2.7 ^{+0.3} _{-0.5}
LMC	NGC 1846	196	81	239.2 ^{+0.2} _{-0.3}	2.1 ^{+0.2} _{-0.4}	269.4 ^{+1.5} _{-1.8}	25.0 ^{+1.3} _{-1.4}
LMC	NGC 1850	155	87	248.9 ^{+0.4} _{-0.5}	2.5 ^{+1.8} _{-0.4}	257.4 ^{+2.8} _{-2.4}	23.6 ^{+1.7} _{-2.6}
LMC	NGC 1978	145	86	293.1 ^{+0.3} _{-0.3}	3.1 ^{+0.3} _{-0.4}	283.7 ^{+3.0} _{-2.2}	25.9 ^{+2.2} _{-2.1}
LMC	NGC 2121	109	72	237.0 ^{+0.3} _{-0.2}	1.8 ^{+0.1} _{-0.4}	262.8 ^{+2.7} _{-3.1}	21.5 ^{+2.0} _{-2.3}
LMC	NGC 2155	110	104	315.0 ^{+0.1} _{-0.3}	0.8 ^{+0.1} _{-0.3}	310.4 ^{+2.1} _{-1.8}	20.7 ^{+1.3} _{-1.6}
LMC	NGC 2203	96	77	252.8 ^{+0.3} _{-0.2}	1.9 ^{+0.2} _{-0.3}	244.2 ^{+4.3} _{-4.8}	23.0 ^{+4.9} _{-5.2}
LMC	NGC 2209	113	112	251.2 ^{+0.1} _{-0.4}	0.8 ^{+1.2} _{-0.2}	251.4 ^{+2.1} _{-1.6}	16.5 ^{+1.3} _{-1.3}
LMC	NGC 2257	94	65	301.8 ^{+0.3} _{-0.4}	2.7 ^{+0.2} _{-0.4}	320.4 ^{+1.3} _{-2.5}	12.8 ^{+1.6} _{-2.0}
LMC	SL 663	113	113	301.1 ^{+1.4} _{-1.2}	3.2 ^{+6.4} _{-2.2}	300.1 ^{+2.2} _{-2.7}	19.5 ^{+2.5} _{-1.6}

4.4 Final kinematic sample

After the addition of 193 more stars from the literature, our final working sample consists of 2787 stars in 26 clusters. We will refer to this as the ‘Kinematic’ sample, and use it exclusively for the final EM analysis following the procedures described in S19 and Section 4.

A summary of the initial analysis of the full kinematic sample is provided in Table 7: Column 3 lists the total number of stars in our kinematic sample for each cluster, while column 4 lists the number of stars within each cluster’s tidal radius (see Table 2); columns 5 and 6 list the estimates of the cluster systemic velocity and projected central velocity dispersion of each cluster; columns 7 and 8 list the mean velocity and velocity dispersion of the field population when applicable.

5 DYNAMICAL AND CHEMICAL RESULTS

5.1 Cluster systemic velocity and velocity dispersion

The systemic velocity (V_{sys}) and projected central velocity dispersion ($\sigma_{p,0}$) listed in Table 7 are the two direct kinematic parameters determined by the EM analysis for each cluster. As discussed in S19, the ‘raw’ EM estimates for these parameters rely critically on the membership probabilities, some of which may be skewed due to field contamination that may result from the relatively small samples being used. Here we follow and expand upon our approach in S19, by calculating V_{sys} and $\sigma_{p,0}$ using two different sampling methods, ‘BOX’ and ‘PM50’. Both use EM results – systemic velocity in the case of BOX, modified membership probabilities in the case of PM50 – to refine the cluster kinematic parameters. The value of

these methods lies in their utility to better understand the effect that marginal outlier stars may have in the final kinematic results for a given cluster.

In the BOX method, V_{sys} and $\sigma_{p,0}$ are computed from all stars in a rectangular area in the $v_{\text{los}}-r$ plane (see the shaded area in fig. 9 of S19). This BOX area is bounded by the cluster centre and the adopted K66 tidal radius (see Table 2) in x -axis, while in y -axis the minimum and maximum systemic velocities are set by conservatively assuming the target cluster has a (large) total mass of $10^6 M_{\odot}$. The value of the BOX method is that it is comparatively easy to specify and compute at the cost of effectively ignoring the EM analysis membership probabilities.

In the PM50 method, V_{sys} and $\sigma_{p,0}$ are computed from probable member stars based on their EM membership probabilities (see Section 4.4). A star is assumed to be a cluster member if its initial EM membership probability is greater than or equal to 50 per cent (i.e. $P_M \geq 0.5$). Unlike the full EM analysis in which membership probabilities are used to weight individual stars, the PM50 method implicitly assigns 100 per cent membership probabilities (therefore assumed to be certain members) to any stars with $P_M \geq 0.5$, and assigns certain non-membership (0 per cent membership probability) to the rest stars in the sample.

Table 8 lists the BOX and PM50 results for all clusters in our sample. We list the numbers of members assigned by each method, along with the corresponding values of V_{sys} and $\sigma_{p,0}$ (columns 3–5 for BOX; columns 6–8 for PM50). The left-hand panel of Fig. 7 compares $\sigma_{p,0}$ estimated from the BOX and PM50 methods. One obvious and unsurprising conclusion from these figures is that the BOX method systematically overestimates the dispersion compared to PM50. Since the BOX method has no provision for flagging

Table 8. Results of the BOX, PM50, and PM50' methods.^a

Galaxy	Field	N_{BOX}	$V_{\text{sys,BOX}}$ (km s ⁻¹)	$\sigma_{\text{p,0,BOX}}$ (km s ⁻¹)	N_{PM50}	$V_{\text{sys,PM50}}$ (km s ⁻¹)	$\sigma_{\text{p,0,PM50}}$ (km s ⁻¹)	$N_{\text{PM50}'}$	$V_{\text{sys,PM50}'}$ (km s ⁻¹)	$\sigma_{\text{p,0,PM50}'}$ (km s ⁻¹)
(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)	(9)	(10)	(11)
SMC	Kron 3	51	132.5 ^{+0.3} _{-0.3}	2.8 ^{+0.3} _{-0.4}	41	132.7 ^{+0.2} _{-0.3}	2.1 ^{+0.2} _{-0.3}	39	132.9 ^{+0.3} _{-0.2}	1.9 ^{+0.2} _{-0.2}
SMC	Lindsay 1	83	140.5 ^{+0.2} _{-0.2}	1.9 ^{+0.2} _{-0.3}	80	140.5 ^{+0.2} _{-0.2}	1.7 ^{+0.2} _{-0.2}	79	140.5 ^{+0.2} _{-0.2}	1.7 ^{+0.2} _{-0.2}
SMC	NGC 152	27	172.8 ^{+0.4} _{-0.4}	2.4 ^{+0.3} _{-0.5}	21	172.6 ^{+0.5} _{-0.5}	2.8 ^{+0.6} _{-1.4}	17	172.6 ^{+0.3} _{-0.3}	1.1 ^{+0.1} _{-0.2}
SMC	NGC 330	18	152.2 ^{+0.9} _{-0.9}	5.1 ^{+0.9} _{-1.4}	14	153.0 ^{+0.6} _{-0.7}	2.4 ^{+0.4} _{-0.8}	13	152.8 ^{+0.6} _{-0.7}	2.3 ^{+0.3} _{-0.9}
SMC	NGC 339	35	112.9 ^{+0.3} _{-0.3}	1.8 ^{+0.2} _{-0.3}	35	112.9 ^{+0.3} _{-0.3}	1.8 ^{+0.2} _{-0.4}	34	113.0 ^{+0.3} _{-0.3}	1.6 ^{+0.1} _{-0.3}
SMC	NGC 361	24	170.4 ^{+0.7} _{-0.7}	5.0 ^{+0.7} _{-0.9}	20	170.5 ^{+0.6} _{-0.6}	3.8 ^{+0.6} _{-0.9}	20	170.9 ^{+0.5} _{-0.5}	3.1 ^{+0.5} _{-0.8}
SMC	NGC 411	35	162.8 ^{+0.7} _{-0.7}	6.8 ^{+1.2} _{-1.3}	22	163.8 ^{+0.3} _{-0.3}	1.7 ^{+0.3} _{-0.4}	20	164.0 ^{+0.2} _{-0.2}	1.4 ^{+0.2} _{-0.3}
SMC	NGC 416	25	155.6 ^{+0.8} _{-0.8}	6.3 ^{+1.3} _{-1.5}	19	155.0 ^{+0.6} _{-0.6}	3.4 ^{+0.4} _{-0.6}	18	154.7 ^{+0.5} _{-0.6}	3.1 ^{+0.3} _{-0.6}
SMC	NGC 419	49	189.9 ^{+0.3} _{-0.3}	3.4 ^{+0.5} _{-0.7}	44	189.9 ^{+0.2} _{-0.2}	2.2 ^{+0.2} _{-0.3}	43	189.8 ^{+0.2} _{-0.2}	2.0 ^{+0.2} _{-0.2}
SMC	NGC 458	19	148.8 ^{+0.7} _{-0.7}	4.8 ^{+0.7} _{-1.0}	14	149.0 ^{+0.7} _{-0.8}	3.5 ^{+0.6} _{-1.3}	10	148.8 ^{+0.4} _{-0.4}	1.6 ^{+0.5} _{-1.0}
LMC	Hodge 4	33	313.0 ^{+0.3} _{-0.3}	1.5 ^{+0.1} _{-0.2}	33	312.8 ^{+0.4} _{-0.3}	1.9 ^{+0.2} _{-0.4}	29	313.1 ^{+0.3} _{-0.3}	1.4 ^{+0.2} _{-0.3}
LMC	NGC 1466	25	202.5 ^{+0.5} _{-0.5}	3.6 ^{+0.4} _{-0.6}	25	202.5 ^{+0.5} _{-0.5}	3.6 ^{+0.4} _{-0.6}	24	202.7 ^{+0.5} _{-0.5}	3.3 ^{+0.4} _{-0.5}
LMC	NGC 1751	46	240.3 ^{+0.5} _{-0.6}	5.1 ^{+0.6} _{-0.9}	28	240.3 ^{+0.4} _{-0.4}	2.5 ^{+0.3} _{-0.6}	21	241.3 ^{+0.2} _{-0.2}	1.1 ^{+0.1} _{-0.2}
LMC	NGC 1783	75	279.5 ^{+0.2} _{-0.2}	2.5 ^{+0.2} _{-0.3}	66	279.6 ^{+0.2} _{-0.2}	2.0 ^{+0.2} _{-0.2}	65	279.5 ^{+0.2} _{-0.2}	1.9 ^{+0.2} _{-0.2}
LMC	NGC 1806	43	230.0 ^{+0.5} _{-0.5}	4.2 ^{+0.6} _{-0.7}	35	229.7 ^{+0.3} _{-0.3}	2.3 ^{+0.3} _{-0.4}	33	230.0 ^{+0.3} _{-0.3}	1.9 ^{+0.2} _{-0.2}
LMC	NGC 1831	77	276.4 ^{+0.3} _{-0.3}	3.7 ^{+0.4} _{-0.5}	64	276.8 ^{+0.2} _{-0.2}	1.9 ^{+0.2} _{-0.2}	62	276.9 ^{+0.2} _{-0.2}	1.7 ^{+0.2} _{-0.2}
LMC	NGC 1841	64	210.8 ^{+0.3} _{-0.3}	2.7 ^{+0.3} _{-0.5}	64	210.8 ^{+0.3} _{-0.3}	2.7 ^{+0.3} _{-0.5}	63	210.6 ^{+0.3} _{-0.3}	2.4 ^{+0.3} _{-0.3}
LMC	NGC 1846	55	239.2 ^{+0.3} _{-0.3}	2.5 ^{+0.3} _{-0.4}	53	239.2 ^{+0.2} _{-0.2}	2.1 ^{+0.2} _{-0.3}	51	239.1 ^{+0.2} _{-0.2}	1.9 ^{+0.2} _{-0.2}
LMC	NGC 1850	74	249.0 ^{+0.4} _{-0.4}	4.3 ^{+0.4} _{-0.5}	63	248.8 ^{+0.3} _{-0.3}	2.6 ^{+0.3} _{-0.3}	62	248.9 ^{+0.3} _{-0.3}	2.5 ^{+0.3} _{-0.3}
LMC	NGC 1978	76	293.0 ^{+0.3} _{-0.3}	3.3 ^{+0.3} _{-0.3}	75	293.1 ^{+0.3} _{-0.3}	3.1 ^{+0.2} _{-0.3}	74	293.2 ^{+0.3} _{-0.3}	2.9 ^{+0.2} _{-0.3}
LMC	NGC 2121	49	236.9 ^{+0.3} _{-0.3}	2.6 ^{+0.4} _{-0.5}	43	237.0 ^{+0.2} _{-0.2}	1.8 ^{+0.2} _{-0.3}	42	236.9 ^{+0.3} _{-0.2}	1.8 ^{+0.2} _{-0.3}
LMC	NGC 2155	57	314.9 ^{+0.4} _{-0.4}	4.4 ^{+0.9} _{-1.0}	35	315.0 ^{+0.2} _{-0.1}	0.8 ^{+0.1} _{-0.3}	34	315.0 ^{+0.2} _{-0.2}	0.8 ^{+0.1} _{-0.3}
LMC	NGC 2203	73	252.9 ^{+0.2} _{-0.2}	1.9 ^{+0.2} _{-0.3}	72	252.8 ^{+0.2} _{-0.2}	1.9 ^{+0.2} _{-0.3}	71	252.9 ^{+0.2} _{-0.2}	1.9 ^{+0.2} _{-0.3}
LMC	NGC 2209	74	250.9 ^{+0.3} _{-0.3}	3.8 ^{+0.5} _{-0.6}	52	251.2 ^{+0.2} _{-0.2}	0.8 ^{+0.2} _{-0.3}	51	251.3 ^{+0.1} _{-0.1}	0.7 ^{+0.1} _{-0.2}
LMC	NGC 2257	63	301.8 ^{+0.3} _{-0.3}	2.7 ^{+0.3} _{-0.3}	63	301.8 ^{+0.3} _{-0.3}	2.7 ^{+0.2} _{-0.3}	61	301.6 ^{+0.3} _{-0.3}	2.4 ^{+0.2} _{-0.3}
LMC	SL 663	32	301.0 ^{+0.4} _{-0.4}	2.2 ^{+0.3} _{-0.4}	23	301.2 ^{+0.5} _{-0.5}	2.6 ^{+0.6} _{-1.1}	20	301.0 ^{+0.3} _{-0.2}	0.9 ^{+0.3} _{-0.6}

Note. ^aAs discussed in Section 5.1, the PM50' results for NGC 458 and SL 663 should be used with caution.

non-members, it is more likely to include random outlier stars whose inclusion depends on how the box areas are (arbitrarily) assigned.

To explore how marginal outlier stars may affect the PM50 method, we have defined a third approach denoted PM50'. For this method, we excluded one PM50 star exhibiting the most deviant velocity (in absolute value) from the PM50 systemic velocity, and re-applied the EM analysis to the remaining full kinematic sample of the cluster. Again, stars are taken to be cluster members if their re-determined EM membership probabilities are greater or equal to 50 per cent, and only such stars are used to compute V_{sys} and $\sigma_{\text{p,0}}$ for the PM50' sample; these results are listed in columns 9–11 of Table 8. Note that the numbers of stars used in the PM50 and PM50' samples of a given cluster do not always differ by one, because when the most probable outlier is removed from the PM50 sample, the EM analysis reassigns membership probabilities that may reassign more than one star as members or non-members.

We compare the PM50 and PM50' results in the right-hand panel of Fig. 7. For most clusters, the change in $\sigma_{\text{p,0}}$ is within 1- σ to the expectation of removing the most deviant star from a Gaussian distribution (grey bands in the right-hand panel of Fig. 7). We conclude that for most cases, the adopted PM50 distributions are consistent with being Gaussian and with the EM membership assignments for individual stars in these distributions.

However, for five clusters – NGC 152, NGC 458, NGC 1751, SL 663 and, marginally, Hodge 4 – we find the PM50 results to be overly sensitive to the removal of a single (outlier) star in the EM analysis. These clusters also exhibit a large change in the number of likely members from their respective PM50 samples as a result of removing one extreme star (see Table 8). For these clusters, we have chosen to adopt their PM50' velocity dispersions from Table 8.

To test the internal consistency of adopting the PM50' results, we conducted the same test as above for the five PM50' clusters. The results are shown in the right-hand panel of Fig. 7. We find that the PM50' samples for three of these clusters – NGC 152, NGC 1751, and Hodge 4 – are now consistent with being Gaussian given their updated EM membership assignments. Two clusters, SL 663 and NGC 458, remain anomalous. Rather than iterate further on these clusters' distributions, we conclude that these two clusters' samples may have issues that affect our EM analyses of the systems. We note that NGC 458 has the smallest stellar sample of all the clusters in our study, and that the structural parameters for SL 663 are particularly uncertain; both factors may lead the EM process astray. This suggests that even the PM50' results for these two clusters may be systematically suspect.

The preferred velocity dispersion and number of cluster members for each cluster – based on either the PM50 or PM50' results – are highlighted with bold font in Table 8.

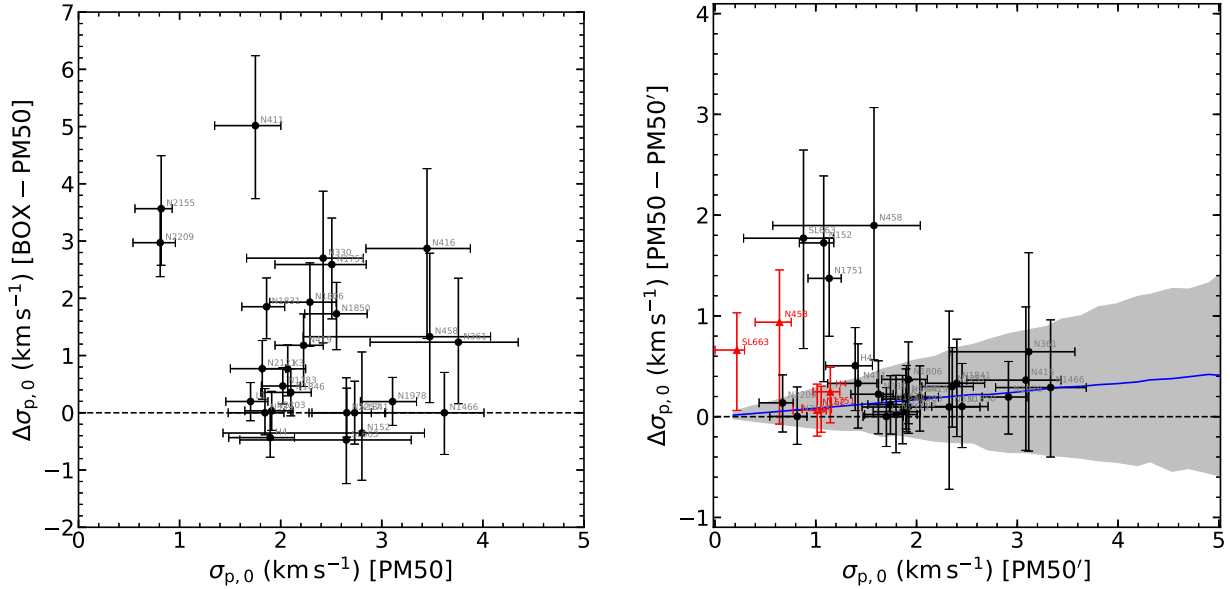


Figure 7. Comparisons of estimated $\sigma_{p,0}$ between the BOX and PM50 methods (left-hand panel), and between the PM50 and PM50' methods (right-hand panel). The $\sigma_{p,0}$ values used in the plots are listed in Table 8. In the right-hand panel, the grey band denotes the 68 per cent ($\pm 1\sigma$) range of 1000 simulated results obtained by removing the most extreme star from the tail of Gaussian distributions consisting of 30 velocities with initial $\sigma_{p,0}$ value ranging between 0.2 and 8.0 km s^{-1} in 0.2 km s^{-1} intervals. The blue line denotes the median results of these simulations. The red triangles denote the results of replicating the process of producing the PM50' sample to the five cases (NGC 152, NGC 458, NGC 1751, SL 663 and Hodge 4) whose black points lie outside the grey band. These results imply that the PM50' samples for NGC 458 and SL 663 are not strictly consistent with Gaussian distributions while the PM50' samples for the other three are consistent. Full details are discussed in Section 5.1.

5.2 Cluster mass and mass-to-light ratio

A cluster's total luminosity and dynamical mass can be derived by scaling the central surface brightness and the projected central velocity dispersion, respectively, with the dimensionless cluster structural profiles (i.e. the K66 models listed in Table 2). The two required scaling parameters, central surface brightness and projected central velocity dispersion, were determined, respectively, from aperture photometry (as described in this section) and from the individual stellar velocities (see Section 5.1).

To determine the central surface brightness ($\Sigma_{V,0}$), we integrated the dimensionless K66 model to a maximal reference radius within which an aperture luminosity was measured from the ground-based CCD photometry. The dimensionless K66 models were computed using the code LIMEPY (Gieles & Zocchi 2015), consistent with how we determine cluster masses (see below). The aperture magnitude and associated reference radius used for each cluster are listed in Table 1, columns 3–4, with the sources identified in column 5. When transforming magnitude to luminosity, we adopted $M_{V,\odot} = 4.85$ in addition to the distance modulus and extinction value listed in Table 1. We could not find published V -band aperture photometry for SL 663, so instead we adopted the central surface brightness reported by McLaughlin & van der Marel (2005). This value, and its uncertainty, is given in Table 1.

The total luminosity ($L_{V,\text{tot}}$) was obtained by integrating the $\Sigma_{V,0}$ scaled K66 model to the K66 tidal radius listed in Table 2. Similarly, the total cluster mass (M_{tot}) was derived with the same dimensionless K66 model but scaled to $\sigma_{p,0}$ fitted from the PM50 or PM50' sample (see Section 5.1 and Table 8). The dynamical M/L_V ratios of the clusters were then derived from the total masses and total V -band luminosity determined above.

Uncertainties in cluster masses, luminosities, and M/L_V ratios were estimated using a bootstrapping technique that accounted for

uncertainties in distance modulus and extinction (Table 1), the K66 structural parameters (see Table 2), the V -band aperture magnitude (for $L_{V,\text{tot}}$ and M/L_V ; see Table 1), and the central velocity dispersion (for M_{tot} and M/L_V ; see Table 8). All relevant parameters were randomly sampled, assuming that the error distributions are joint-Gaussian to account for the asymmetric uncertainties (e.g. c_{K66} in Table 2), or using simple Gaussian distributions for parameters with symmetric uncertainties (e.g. V_{ap} in Table 1). For each cluster, we created 1000 samples from the selected parameters to calculate M_{tot} , $L_{V,\text{tot}}$, and M/L_V . We took the 15.9-th and 84.1-th percentiles of the simulated values as the lower and upper 1σ confidence boundaries of the respective parameter. Table 9 lists the masses, luminosities, and M/L_V ratios for all 26 clusters in our sample.

Our dynamical analysis assumed that all clusters are exclusively pressure-supported systems. However, for two cases – NGC 1978 and NGC 1846 – the kinematic data suggest that the clusters may exhibit coherent rotation (see S19 and Fig. 8). For NGC 1846, we roughly addressed the dynamical effect of rotation in S19, concluding that rotation may be causing ~ 9 per cent overestimation in the cluster's total mass and M/L_V . For NGC 1978, Fischer et al. (1992b) carried out a more sophisticated analysis in which they fitted single-mass rotating and non-rotating oblate spheroid models to the surface luminosity profiles and their radial velocity data. They found no significant differences in M/L_V s derived with these models and those derived with single-mass K66 models. However, the mass estimates for the cluster did differ systematically among the rotating and non-rotating models. Our data for NGC 1978, while clearly showing a rotation signature (see right-hand panel of Fig. 8), are limited by poor background determination as noted in Section 2.4.2. For this reason, and because NGC 1978 requires a more involved dynamical modelling approach, we will defer a detailed analysis for this cluster to a later paper.

Table 9. Mass, luminosity, and M/L_V .^a

Galaxy	Cluster	Method	M_{tot} ($\times 10^5 M_{\odot}$)	$L_{V, \text{tot}}$ ($\times 10^5 L_{\odot}$)	M/L_V ($M_{\odot} L_{\odot}^{-1}$)	$\log M_{\text{tot}}$ (M_{\odot})	$\log L_{V, \text{tot}}$ (L_{\odot})	$\log M/L_V$ ($M_{\odot} L_{\odot}^{-1}$)
(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)	(9)
SMC	Kron 3	PM50	$0.77^{+0.15}_{-0.20}$	$1.41^{+0.18}_{-0.14}$	$0.55^{+0.11}_{-0.14}$	$4.89^{+0.08}_{-0.13}$	$5.15^{+0.05}_{-0.05}$	$-0.26^{+0.08}_{-0.13}$
SMC	Lindsay 1	PM50	$0.76^{+0.16}_{-0.20}$	$0.87^{+0.09}_{-0.08}$	$0.88^{+0.20}_{-0.24}$	$4.88^{+0.08}_{-0.13}$	$4.94^{+0.04}_{-0.04}$	$-0.05^{+0.09}_{-0.14}$
SMC	NGC 152	PM50	$1.67^{+0.83}_{-1.27}$	$0.80^{+0.12}_{-0.10}$	$2.09^{+0.94}_{-1.60}$	$5.22^{+0.18}_{-0.62}$	$4.90^{+0.06}_{-0.06}$	$0.32^{+0.16}_{-0.63}$
		PM50'	$0.25^{+0.07}_{-0.10}$	$0.80^{+0.13}_{-0.09}$	$0.31^{+0.06}_{-0.13}$	$4.39^{+0.10}_{-0.22}$	$4.90^{+0.07}_{-0.06}$	$-0.51^{+0.08}_{-0.23}$
SMC	NGC 330	PM50	$0.54^{+0.22}_{-0.30}$	$8.93^{+1.05}_{-0.86}$	$0.06^{+0.02}_{-0.03}$	$4.74^{+0.15}_{-0.34}$	$5.95^{+0.05}_{-0.04}$	$-1.22^{+0.13}_{-0.33}$
SMC	NGC 339	PM50	$0.57^{+0.23}_{-0.23}$	$0.72^{+0.28}_{-0.16}$	$0.79^{+0.15}_{-0.29}$	$4.76^{+0.15}_{-0.23}$	$4.86^{+0.14}_{-0.10}$	$-0.10^{+0.07}_{-0.20}$
SMC	NGC 361	PM50	$2.15^{+1.36}_{-0.98}$	$1.04^{+0.54}_{-0.31}$	$2.07^{+0.75}_{-0.85}$	$5.33^{+0.21}_{-0.26}$	$5.02^{+0.18}_{-0.15}$	$0.32^{+0.13}_{-0.23}$
SMC	NGC 411	PM50	$0.30^{+0.10}_{-0.13}$	$0.80^{+0.10}_{-0.08}$	$0.38^{+0.11}_{-0.16}$	$4.48^{+0.12}_{-0.24}$	$4.90^{+0.05}_{-0.04}$	$-0.42^{+0.11}_{-0.23}$
SMC	NGC 416	PM50	$0.80^{+0.24}_{-0.30}$	$1.12^{+0.13}_{-0.10}$	$0.72^{+0.17}_{-0.25}$	$4.90^{+0.11}_{-0.20}$	$5.05^{+0.05}_{-0.04}$	$-0.15^{+0.09}_{-0.19}$
SMC	NGC 419	PM50	$0.64^{+0.14}_{-0.15}$	$3.49^{+0.62}_{-0.51}$	$0.18^{+0.05}_{-0.05}$	$4.80^{+0.09}_{-0.12}$	$5.54^{+0.07}_{-0.07}$	$-0.74^{+0.10}_{-0.13}$
SMC	NGC 458	PM50	$1.26^{+0.58}_{-0.73}$	$1.24^{+0.36}_{-0.26}$	$1.02^{+0.44}_{-0.62}$	$5.10^{+0.16}_{-0.37}$	$5.09^{+0.11}_{-0.10}$	$0.01^{+0.16}_{-0.40}$
		PM50'	$0.26^{+0.21}_{-0.22}$	$1.24^{+0.35}_{-0.24}$	$0.21^{+0.15}_{-0.18}$	$4.42^{+0.26}_{-0.80}$	$5.09^{+0.11}_{-0.09}$	$-0.68^{+0.23}_{-0.80}$
LMC	Hodge 4	PM50	$3.39^{+229.01}_{-2.49}$	$2.67^{+202.43}_{-1.82}$	$1.27^{+0.28}_{-0.54}$	$5.53^{+1.84}_{-0.58}$	$5.43^{+1.89}_{-0.50}$	$0.10^{+0.09}_{-0.24}$
		PM50'	$1.82^{+125.83}_{-1.33}$	$2.67^{+228.50}_{-1.81}$	$0.68^{+0.17}_{-0.27}$	$5.26^{+1.84}_{-0.57}$	$5.43^{+1.94}_{-0.49}$	$-0.17^{+0.09}_{-0.22}$
LMC	NGC 1466	PM50	$1.02^{+0.27}_{-0.31}$	$0.94^{+0.09}_{-0.07}$	$1.09^{+0.25}_{-0.33}$	$5.01^{+0.10}_{-0.16}$	$4.97^{+0.04}_{-0.03}$	$0.04^{+0.09}_{-0.15}$
LMC	NGC 1751	PM50	$0.76^{+0.23}_{-0.31}$	$0.90^{+0.14}_{-0.12}$	$0.84^{+0.27}_{-0.34}$	$4.88^{+0.11}_{-0.23}$	$4.96^{+0.06}_{-0.06}$	$-0.08^{+0.12}_{-0.23}$
		PM50'	$0.15^{+0.03}_{-0.05}$	$0.90^{+0.13}_{-0.11}$	$0.17^{+0.04}_{-0.06}$	$4.19^{+0.09}_{-0.19}$	$4.96^{+0.06}_{-0.06}$	$-0.77^{+0.10}_{-0.19}$
LMC	NGC 1783	PM50	$0.98^{+0.17}_{-0.20}$	$3.77^{+0.28}_{-0.27}$	$0.26^{+0.04}_{-0.05}$	$4.99^{+0.07}_{-0.10}$	$5.58^{+0.03}_{-0.03}$	$-0.58^{+0.07}_{-0.09}$
LMC	NGC 1806	PM50	$0.76^{+0.18}_{-0.26}$	$1.42^{+0.10}_{-0.10}$	$0.54^{+0.13}_{-0.18}$	$4.88^{+0.09}_{-0.18}$	$5.15^{+0.03}_{-0.03}$	$-0.27^{+0.09}_{-0.18}$
LMC	NGC 1831	PM50	$0.41^{+0.09}_{-0.11}$	$1.33^{+0.18}_{-0.17}$	$0.31^{+0.07}_{-0.08}$	$4.61^{+0.09}_{-0.14}$	$5.12^{+0.06}_{-0.06}$	$-0.51^{+0.09}_{-0.14}$
LMC	NGC 1841	PM50	$1.04^{+2.53}_{-0.25}$	$0.74^{+1.63}_{-0.02}$	$1.40^{+0.49}_{-0.51}$	$5.02^{+0.54}_{-0.12}$	$4.87^{+0.50}_{-0.01}$	$0.14^{+0.13}_{-0.20}$
LMC	NGC 1846	PM50	$0.57^{+0.13}_{-0.14}$	$1.68^{+0.35}_{-0.30}$	$0.34^{+0.11}_{-0.10}$	$4.75^{+0.09}_{-0.12}$	$5.23^{+0.08}_{-0.09}$	$-0.47^{+0.13}_{-0.15}$
LMC	NGC 1850	PM50	$0.52^{+0.05}_{-0.17}$	$6.56^{+1.30}_{-1.07}$	$0.08^{+0.01}_{-0.03}$	$4.71^{+0.04}_{-0.18}$	$5.82^{+0.08}_{-0.08}$	$-1.10^{+0.06}_{-0.20}$
LMC	NGC 1978	PM50	$1.36^{+0.24}_{-0.29}$	$3.41^{+0.30}_{-0.25}$	$0.40^{+0.06}_{-0.08}$	$5.13^{+0.07}_{-0.11}$	$5.53^{+0.04}_{-0.03}$	$-0.40^{+0.06}_{-0.10}$
LMC	NGC 2121	PM50	$0.50^{+0.22}_{-0.19}$	$0.79^{+0.42}_{-0.26}$	$0.63^{+0.20}_{-0.21}$	$4.69^{+0.16}_{-0.21}$	$4.90^{+0.19}_{-0.17}$	$-0.20^{+0.12}_{-0.18}$
LMC	NGC 2155	PM50	$0.06^{+0.03}_{-0.04}$	$0.22^{+0.14}_{-0.08}$	$0.29^{+0.17}_{-0.18}$	$3.81^{+0.16}_{-0.36}$	$4.35^{+0.22}_{-0.18}$	$-0.54^{+0.20}_{-0.42}$
LMC	NGC 2203	PM50	$0.51^{+0.13}_{-0.13}$	$0.81^{+0.12}_{-0.11}$	$0.63^{+0.19}_{-0.18}$	$4.70^{+0.10}_{-0.13}$	$4.91^{+0.06}_{-0.06}$	$-0.20^{+0.11}_{-0.14}$
LMC	NGC 2209	PM50	$0.09^{+0.02}_{-0.06}$	$0.26^{+0.02}_{-0.02}$	$0.36^{+0.08}_{-0.21}$	$3.97^{+0.09}_{-0.40}$	$4.41^{+0.03}_{-0.03}$	$-0.45^{+0.09}_{-0.39}$
LMC	NGC 2257	PM50	$1.01^{+0.68}_{-0.38}$	$0.51^{+0.39}_{-0.17}$	$2.00^{+0.38}_{-0.53}$	$5.01^{+0.22}_{-0.20}$	$4.71^{+0.25}_{-0.17}$	$0.30^{+0.08}_{-0.13}$
LMC	SL 663	PM50	$3.32^{+778.97}_{-2.65}$	$0.45^{+132.55}_{-0.33}$	$7.31^{+5.34}_{-4.81}$	$5.52^{+2.37}_{-0.70}$	$4.66^{+2.47}_{-0.55}$	$0.86^{+0.24}_{-0.47}$
		PM50'	$0.36^{+69.09}_{-0.33}$	$0.45^{+164.23}_{-0.36}$	$0.80^{+0.71}_{-0.70}$	$4.56^{+2.28}_{-1.03}$	$4.66^{+2.56}_{-0.67}$	$-0.10^{+0.27}_{-0.87}$

Note. ^aAs discussed in Section 5.1, the PM50' results for NGC 458 and SL 663 should be used with caution.

5.3 Comparison with previous work

Previously published studies have reported velocity dispersions and M/L ratios of 13 clusters in common with our work. Seven of these clusters were studied using integrated-light spectroscopy (Dubath, Meylan & Mayor 1997; Zaritsky et al. 2012, 2014), and another six clusters were studied from dynamical modelling using radial velocities of individual cluster member stars similar to this work (Fischer et al. 1992b, 1993; Suntzeff et al. 1992; Mackey et al. 2013; Kamann et al. 2018b; Patrick et al. 2020).

The left-hand panel of Fig. 9 shows the difference in central velocity dispersions between these previous studies and our work (see Tables 8 and 9). Our dispersion estimates agree well with studies using individual stellar spectra (red triangles); the average difference (red dashed line) is $0.37 \pm 0.19 \text{ km s}^{-1}$. As for the integrated-light studies (blue and green squares), only two clusters with relatively high central values (i.e. NGC 419 and NGC 1466) agree in the

dispersions, while we measured lower values for the remaining five clusters. For those studied by Zaritsky et al. (2012), Zaritsky et al. (2014) (blue squares), we obtained an average difference of $1.13 \pm 0.20 \text{ km s}^{-1}$ in dispersion (blue dashed line).

In the right-hand panel of Fig. 9, we show the difference in M/L_V ratios between these previous studies and our work (see Tables 8 and 9). In general, our measured M/L_V ratios agree better with those studied by Zaritsky et al. (2012, 2014), compared to those using individual stellar spectra. For the clusters studied by Zaritsky et al. (2012, 2014) (blue squares), we obtain a weighted average of 0.14 ± 0.08 in $\Delta \log_{10}(M/L_V)$ ratio (blue dashed curves); while for the clusters studied using individual stellar spectra (red triangles), the weighted average (red dashed curves) is 0.23 ± 0.08 . These differences may be rooted from the various modelling methods used to estimate cluster M/L_V ratio. Indeed, our results are in good agreement with those of the studies

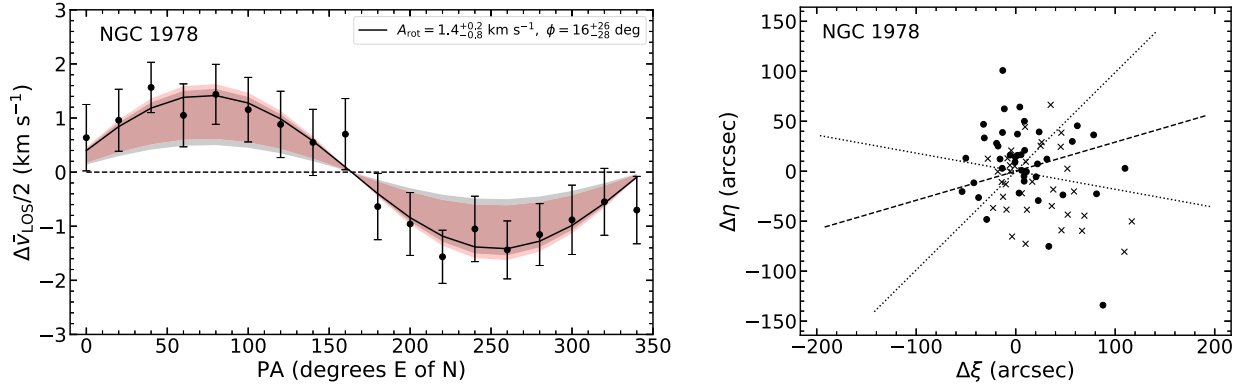


Figure 8. Simple rotation analysis for the stars with $P_M \geq 0.5$ in NGC 1978 (see Section 5.2) using the cluster centres listed in Table 2. In the left-hand panel, we show $\Delta \bar{V}_{\text{los}}/2$ as a function of the bisector position angle (PA), together with the best-fitting sinusoidal model $\Delta \bar{V}_{\text{los}}/2 = A_{\text{rot}} \sin(\text{PA} + \phi)$ (see S19 for details). The best-fitting parameters are listed in the figure legend. In the right-hand panel, crosses (dots) indicate stars with velocities greater (less) than the systemic velocity. The best-fitting rotation axis from the left-hand panel is marked as a dash line in the right-hand panel, and two dotted lines denote the 1σ uncertainties. Apart from NGC 1978 shown in this figure, the rotation signature of all other clusters in our sample is negligible ($A_{\text{rot}}/\sigma_{p,0} \lesssim 0.3$).

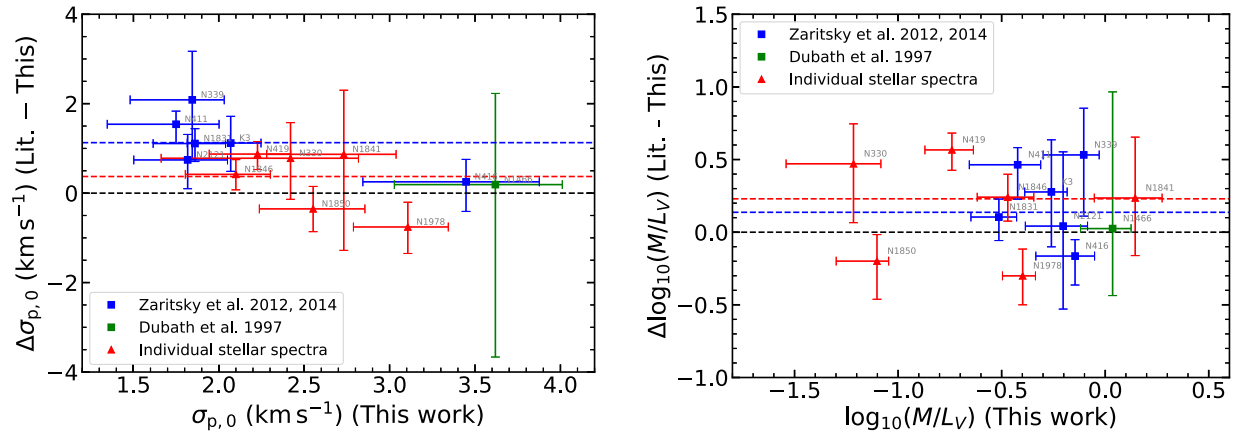


Figure 9. Differences between dynamical results – central velocity dispersions (left-hand panel) and M/L_V ratios (right-hand panel) – found in previous studies compared to results of this paper. In both panels, different symbols indicate different spectroscopic measurements used in those studies (see Section 5.3). The blue and dashed horizontal lines denote the weighted-average offset values of the corresponding coloured data points.

which adopted the K66 models (i.e. NGC 1466, NGC 1841, and NGC 1850).⁷

Overall, the variations in M/L_V ratios between our work and previous studies seem reasonable given difference in observational technique, data quality, sample size, and analysis methods. This conclusion is consistent with the detailed case study for two clusters – NGC 419 and NGC 1846 – in S19, where we compared our results with similar studies based on individual stellar velocity measurements (Mackey et al. 2013; Kamann et al. 2018b). We explored the effects of both poorly estimated velocity uncertainties (for NGC 419) and different dynamical modelling (for NGC 1846) on the determination of M/L_V ratios. We presume that similar issues affect the comparisons of results described here, though for the present sample, and in particular for the integrated-light spectroscopy

⁷This discussion should not be construed to imply broad agreement between our work and that of Zaritsky et al. (2012, 2014). As it happens, the only clusters in common between these studies have intermediate ages. As we shall see in Section 6.4, our M/L estimates do not agree as well for younger and older clusters where our results are systematically lower and higher, respectively.

results (Dubath et al. 1997; Zaritsky et al. 2012, 2014), we are unable to carry out as detailed a comparison as in S19.

5.4 Cluster metallicity

As described in Section 3.2, we estimated effective temperature for each stellar target from their colours in order to break the strong temperature–metallicity degeneracy in our Bayesian fitting analysis (see Section 3.1). The stellar metallicities obtained in this analysis were used in Section 4.2 to flag chemically discrepant stars as non-members before estimating the ‘raw’ mean metallicities of the clusters.

An open question remains as to how well these ‘raw’ metallicities compare to previous chemical analyses. One extensive and consistent source of cluster metallicities come from Ca II triplet (CaT) results that are typically calibrated to two [Fe/H] scales, i.e. Zinn & West (1984) (ZW84) and Carretta & Gratton (1997) (CG97). For our target clusters, we have collected all the studies using either [Fe/H] scale, and then used the relation in Carretta et al. (2001, see their equation 3) to transfer any ZW84 [Fe/H] values on to the CG97 scale. In the left-hand panel of Fig. 10, we compared the [Fe/H] values of nine LMC clusters from Grocholski et al. (2006) (crosses) and five

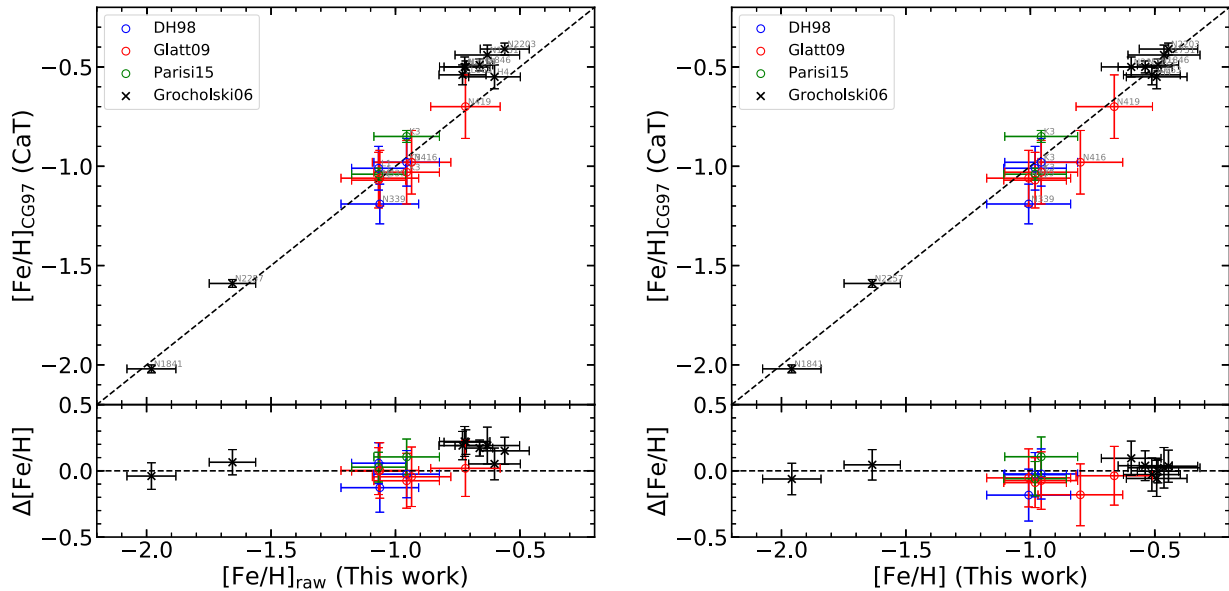


Figure 10. Comparisons of metallicities for clusters is common between our work and that of previous CaT studies (Da Costa & Hatzidimitriou 1998; Grocholski et al. 2006; Glatt et al. 2009; Parisi et al. 2015). In both panels, open circles denote the SMC clusters and crosses denote the LMC clusters of our sample (see Table 6). The raw cluster metallicities derived in Section 4.2 are used in the left-hand panel, while right-hand panel uses the cluster metallicities (as listed in Table 6) obtained after the calibration described in Section 5.4.

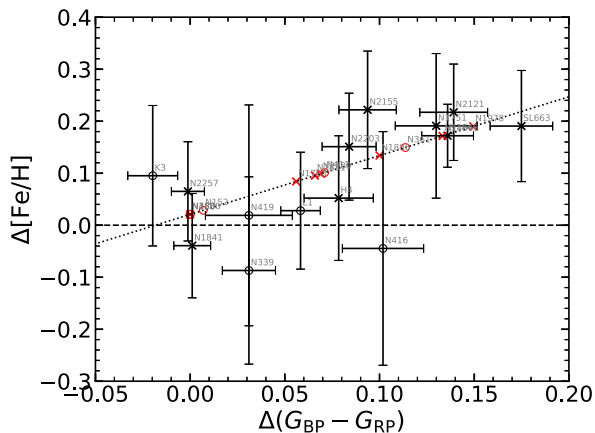


Figure 11. Metallicity differences for clusters in common with our work and previous CaT studies as a function of the colour offsets raised in the T_{eff} calculation (see Section 3.2). The open circles denote the SMC clusters and crosses denote the LMC clusters in our sample (see Table 6). The dotted line is the weighted linear fit to these data given by equation (3). This fitted line was used to place our cluster metallicities on the CaT system described in Section 5.4 and Fig. 10. For clusters in our sample lacking published CaT measurements, the adopted $\Delta[\text{Fe}/\text{H}]$ values are denoted as red symbols along the dotted line. The calibrated cluster metallicities are listed in Table 6.

SMC clusters from Da Costa & Hatzidimitriou (1998), Glatt et al. (2009), Parisi et al. (2015) (open circles). For metal-poor clusters, our raw $[\text{Fe}/\text{H}]$ values agree well with the CaT results; for metal-rich (and mostly LMC) clusters, our raw $[\text{Fe}/\text{H}]$ values are systematically lower than those measured by CaT spectroscopy.

The systematic offset in metallicity is found to be related to the colour offset between the colours used to calculate T_{eff} (see Section 3.2) and those published in the *Gaia* DR2. Fig. 11 shows the metallicity offsets between the CaT studies and our work as a function of the $G_{\text{BP}} - G_{\text{RP}}$ colour offsets of the target clusters in

common. The colour offsets were calculated the same as that shown in Fig. 4 but in a cluster-by-cluster manner, and in the calculation we only considered the cluster members confirmed in Section 4.2. For the three SMC clusters (Kron 3, Lindsay 1, and NGC 339) that have multiple CaT measurements available, we took the weighted average of their $[\text{Fe}/\text{H}]$ values when calculating the metallicity offsets. We found a clear positive correlation between metallicity offsets and colour offsets. We fitted the following linear relation (see the dotted line in Fig. 10) to calibrate our raw cluster metallicities on to the CaT CG97 scale:

$$[\text{Fe}/\text{H}]_{\text{CG97}} = [\text{Fe}/\text{H}]_{\text{raw}} + 1.131\Delta(G_{\text{BP}} - G_{\text{RP}}) + 0.021. \quad (3)$$

The final $[\text{Fe}/\text{H}]_{\text{CG97}}$ values are listed in column 3 of Table 6. The weighted standard deviation in $[\text{Fe}/\text{H}]_{\text{CG97}}$ about the fitted line is 0.064 dex; we have added this value to the final metallicity errors in quadrature (column 4 of Table 6). The comparison between our calibrated cluster metallicities and those in the CG97 scale is shown in the right-hand panel of Fig. 10. The plot shows that along the full metallicity range, our modified $[\text{Fe}/\text{H}]$ values now agree well with those in the CG97 scale after the calibration using equation (3).

6 DISCUSSION

Having produced M/L_V values for 26 clusters spanning a range of ages (Table 1), metallicities (Table 6), and masses (Table 9), we now turn to a discussion regarding the basic trends of M/L_V with respect to these various parameters. Part of that discussion involves a comparison of our results to M/L_V values predicted by a reference set of Simple Stellar Population (SSP) models. Dynamical effects can also lead to evolution of M/L_V values in clusters, and we consider these effects as well. Given that the SSP+dynamics models appear to explain our results reasonably well, we also explore how cluster M/L_V values may be used to constrain other astrophysical parameters such as the stellar IMF and cluster disruption time-scales in the Magellanic Clouds.

6.1 M/L_V trends and comparing to SSP model

A fundamental prediction of stellar evolutionary models is that, barring any strong IMF variations or pathological dynamical effects, simple populations should become ‘darker’ (i.e. higher M/L_V) with increasing age. The top left-hand panel of Fig. 12 reveals such a trend as a clear positive correlation between M/L_V ratio and age for our cluster sample.

In the same panel, we have overplotted evolutionary tracks of a set of SSP models denoted by the solid and dashed black lines and the shaded grey area bounded by these two lines. The tracks were produced using the Flexible Stellar Population Synthesis (FSPS) code (Conroy et al. 2009; Conroy & Gunn 2010) employing Padova isochrones (Girardi et al. 2000; Marigo et al. 2008), a Kroupa (2001) IMF, and the BaSeL spectral library (Lejeune, Cuisinier & Buser 1997, 1998; Westera et al. 2002). We will refer to the tracks produced by this combination of Kroupa/Padova/BaSeL models as the ‘reference’ SSP models throughout this paper. The shaded area in the figure is meant to roughly represent the age–metallicity regime appropriate for MC clusters. The colours of the data points denote age bins used to categorize the clusters. The coloured dotted lines represent the approximate average ages of these bins (see Table 10 and the caption of Fig. 12 for details regarding these age bins).

We also tested other options in the FSPS code, such as using the MILES spectral library (Sánchez-Blázquez et al. 2006) with both MIST (Paxton et al. 2011, 2013, 2015; Choi et al. 2016; Dotter 2016) and BaSTI theoretical isochrones (Pietrinferni et al. 2004). In all cases, we found good agreement ($\lesssim 10$ per cent) with the adopted Padova/Kroupa/BaSeL models among predicted M/L_V ratios using alternative isochrones and stellar libraries.

Fig. 12 reveals that when compared to the adopted reference SSP models, our dynamical M/L_V ratios tend to run lower than the predictions. This offset is highlighted in the top right-hand panel of Fig. 12, where we show the difference between our measured M/L_V ratios and the SSP predictions as a function of cluster age. The $\Delta \log_{10}(M/L_V)$ values were calculated uniquely for each cluster using its own age (see Table 1) and metallicity (see Table 6). Across the entire sample, our M/L_V ratios are -0.24 ± 0.03 dex lower on average than the theoretical predictions in log space, with an error-weighted standard deviation of 0.16 dex.⁸ There is a weak indication that the youngest clusters may more closely follow the SSP predictions. Notably, the overall offset – about 70 per cent – cannot be accounted for by using different input models in the FSPS code.

As shown in the middle left-hand panel of Fig. 12, we also find a trend of decreasing M/L_V ratio with increasing cluster metallicity. This trend is not unexpected since the more metal-rich clusters in our sample are also younger and hence lower in M/L_V ratio than those of the metal-poor counterparts; this is simply a manifestation of the well-known age–metallicity relation for MC clusters (see Fig. 1 and, e.g. Harris & Zaritsky 2009; Parisi et al. 2015). In the middle right-hand panel, the $\Delta \log_{10}(M/L_V)$ values are plotted against the metallicities for all clusters in our sample.

The bottom left-hand panel of Fig. 12 shows a broad trend of increasing M/L_V ratio with increasing cluster mass. This behaviour is also seen among old (globular) clusters in both our Galaxy

(Mandushev et al. 1991; Kimmig et al. 2015) and M31 (Strader et al. 2011), and were attributed by these authors to be due to dynamical evolutionary effects. The isochrones from the reference SSP models (coloured dotted lines) remain constant with cluster mass and so do not predict a trend of M/L_V with total mass. We will return to this when we consider dynamical evolution effects in the MC clusters. The bottom right-hand panel shows the offsets in $\Delta \log_{10}(M/L_V)$ with the SSP predictions calculated using the appropriate age and metallicity for each cluster.

6.2 Dynamical effects on cluster M/L_V values

6.2.1 Mass underestimates from single-mass models

Our dynamical analysis employs single-mass models – specifically K66 models – that assume equal-mass particles and a constant cluster M/L_V ratio at all radii. Such models do not account for dynamical effects associated with energy equipartition that lead to observable features such as mass segregation and, hence, spatial evolution of M/L_V over time. Due to these effects, we can expect that any single-mass models, such as the K66 models we used, will tend to underestimate the total mass, especially when the kinematic tracers (such as our observed RGs) are more massive than the mean mass of cluster members and therefore kinematically colder and more concentrated in the inner regions of the cluster.

Sollima et al. (2015) explored this effect quantitatively by comparing different analytic models (including the K66 model) used to simulated observations obtained from a suite of N -body simulations of star clusters in different stages of their evolution. For clusters with high degree of relaxation (i.e. half-mass relaxation time-scale $t_{rh} = 0.12$ Gyr), they found that the cluster mass can be underestimated up to 50 per cent of the true value. This accounts for a correction of 0.30 dex in $\log_{10}(M/L_V)$ and could, in principle, fully explain the offset seen in Section 6.1. However, it is unlikely that most of our clusters are highly relaxed since some are relatively young and most are comparatively low-density systems.

A more applicable estimation using a longer relaxation time-scale (i.e. $t_{rh} = 4.97$ Gyr) was also studied by Sollima et al. (2015). For this case, they found that K66 models can underestimate the true mass of about 10–20 per cent, depending on the initial cluster mass, the strength of the tidal field, and the radial extent of the kinematics tracers (the RGs in our case) used in the dynamical analysis. This is consistent with the work by Hénault-Brunet et al. (2019), who used mock data from a star-by-star N -body simulation of M4 to compare mass modelling techniques, including ones using K66 models. They found that the K66 model underestimates the true cluster mass by about 17 per cent. This corresponds to an offset in $\log_{10}(M/L_V)$ of 0.08 dex, insufficient to fully account for the offset between our observed M/L_V and the reference SSP models (though it can reduce the offset to about half of what is observed). We conclude that our adoption of K66 models tends to underestimate the true cluster mass, but this does not by itself account for the offset we observe between our measured M/L_V ratios and the reference SSP models (top right-hand panel of Fig. 12).

6.2.2 External dynamical effects

After a bound star cluster forms from a dense gas cloud and survives the so-called ‘infant mortality’ stage (the time-scale of about 10 Myr for unbound clusters to totally dissolve, e.g. Lada & Lada 2003; Whitmore 2004), its evolution will be driven by both internal and external dynamical effects. The internal effects include those

⁸When excluding the two clusters with anomalous PM50 results – NGC 458 and SL 663 (see Section 5.1), the offset in $\Delta \log_{10}(M/L_V)$ of the remaining cluster sample is -0.25 ± 0.03 dex, with an error-weighted standard deviation of 0.16 dex. This shows that the two clusters do not influence the results of our entire cluster sample. The same conclusion applies to all the following statistical results presented in Section 6.

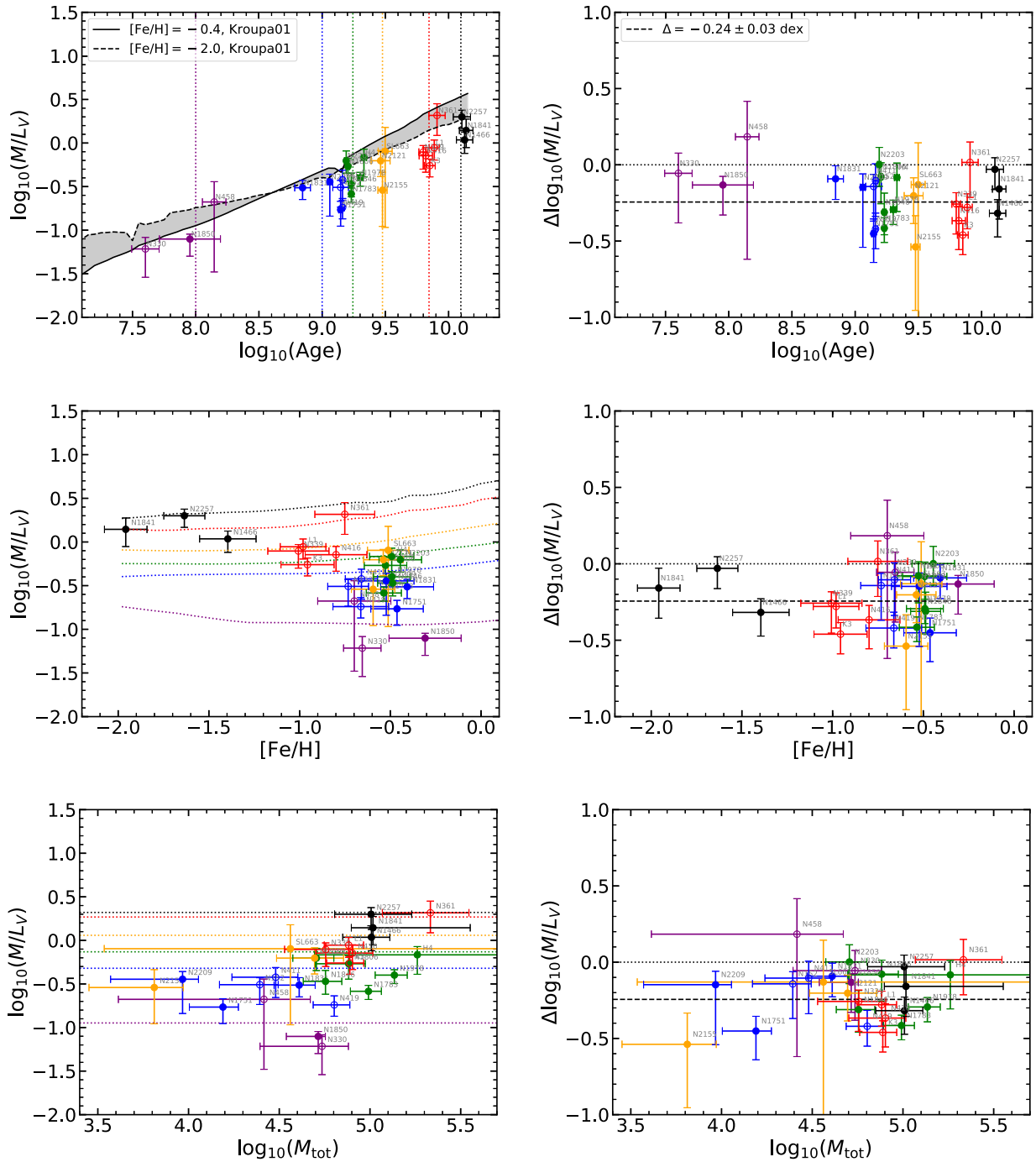


Figure 12. Left-hand panels: Dynamical M/L_V ratios of our target clusters in the LMC (filled circles) and SMC (open circles) as a function of age (top), metallicity (middle), and mass (bottom), overplotted with evolutionary M/L_V ratio isochrones from a set of FSPS SSP models (see Section 6.1 for details). The clusters are binned into distinct age groups denoted by different colours (see Table 10 for bin details). In the top left-hand panel, the solid and dashed curves show M/L_V evolutionary tracks for two metallicities (see the legend). In the top and middle left-hand panels, the coloured dotted lines denote the ages used to generate the SSP isochrone curves (see column 4 of Table 10). In the bottom left-hand panel, the coloured dotted lines are calculated for the same ages but in each case we adopted the bin metallicity of the set of clusters denoted by the same colour (see column 5 of Table 10). Right-hand panels: Differences of M/L_V ratios in log space between our measurements and the SSP predictions. The $\Delta \log_{10}(M/L_V)$ values were calculated in a cluster-by-cluster manner using the age and metallicity of each cluster (we used the ages listed in Table 1 and metallicities listed in Table 6). In each right-hand panel, the horizontal dashed line shows the weighted mean offset in $\Delta \log_{10}(M/L_V)$ for the entire sample.

described in Section 6.2.1 as well as changes in the masses of stars due to mass-loss or binary mergers (e.g. Portegies Zwart, McMillan & Gieles 2010; Renaud 2018). External effects can include tidal perturbations from the host galaxy due to either impulsive effects

– e.g. encounters with giant molecular clouds, spiral arms, or other clusters – or secular evolution arising from a changing tidal field as a cluster orbits within a galaxy (see e.g. Krumholz, McKee & Bland-Hawthorn 2019). In the reference SSP models used here, generally

Table 10. Age/metallicity bins of the cluster sample.

Bin	Colour ^a	Age range (Gyr)	Bin age ^b (Gyr)	Bin [Fe/H] ^b (dex)	Clusters
(1)	(2)	(3)	(4)	(5)	(6)
1	Purple	0.04–0.14	0.1	−0.4	NGC 330, 458, 1850
2	Blue	0.7–1.5	1.0	−0.4	NGC 152, 411, 419, 1751, 1831, 2209
3	Green	1.5–2.2	1.8	−0.4	NGC 1783, 1806, 1846, 1978, 2203; Hodge 4
4	Yellow	2.9–3.2	3.0	−0.4	NGC 2121, 2155; SL 663
5	Red	6.3–8.1	7.0	−0.7	NGC 339, 361, 416; Kron 3, Lindsay 1
6	Black	12.7–13.8	12.5	−1.7	NGC 1466, 1841, 2257

Note. ^aThese colours are used in Figs 12–15 to represent clusters and model results associated with the age/metallicity bins listed here. ^bThese bin age and metallicity values are used to produce the isochrone curves shown in Fig. 12, Fig. 13, and Fig. 14. They are set to equal grid point values in the Anders, Lamers & Baumgardt (2009) models.

only the internal effects related to stellar evolution are considered. Ignoring cluster dynamical evolution – both the internal dynamics and the external tidal effects – can lead to overestimated M/L_V values over time (we have seen this to be the case for internal dynamical evolution in Section 6.2.1).

To test how external dynamical effects may influence cluster M/L_V ratios, we have adapted a suite of evolutionary synthesis models developed by Anders et al. (2009). Their models are built from the GALEV code (see e.g. Kotulla et al. 2009) using the Padova isochrones, a Kroupa (2001) IMF and BaSeL spectral library (the same as the SSP models we adopted in Section 6.1). These models account for dynamical evolution of star cluster by introducing a mass-dependent parameter – the total cluster disruption time, $t_{95 \text{ per cent}}$, defined as the time when 95 per cent of the initial cluster mass is unbound. This time-scale attempts to parametrize the mass-function evolution found in N -body simulations of stars clusters dissolving in tidal fields (Baumgardt & Makino 2003).

In the top left-hand panel of Fig. 13, we compare our empirical M/L_V ratios as a function of age with the Anders et al. (2009) models. For purposes of comparison, we present the Anders et al. (2009) models as coloured bands denoting disruption times set to twice the mean ages of each age group (see the figure caption and Table 10 for details). The models reveal that external dynamical evolution causes the M/L_V ratio to increase slowly then decrease over time as a cluster ages. This is in contrast to the SSP predictions where cluster M/L_V values steadily increase with time (the light grey band in the top left-hand panel of Fig. 13 and, more clearly, in the top left-hand panel of Fig. 12).

The remaining panels of Fig. 13 plot $\Delta \log_{10}(M/L_V)$ as a function of age in a cluster-by-cluster manner. The $\Delta \log_{10}(M/L_V)$ in each panel are determined against the Anders et al. (2009) models assuming different disruption times in an increasing trend, i.e. the disruption time is assumed to be twice the age of a given cluster (top right-hand panel), the sum of the cluster age and its current core relaxation time (bottom left-hand panel), and the sum of the cluster age and its current half-mass relaxation time (bottom right-hand panel), respectively. The core and half-mass relaxation times of each cluster were derived by using the equations listed in Mackey et al. (2013), except that we adopted the Coulomb logarithm as $\Lambda \sim 0.11N$ with N being the total number of stars in the cluster (Giersz & Heggie 1994). All three panels show that our measured M/L_V values agree better with models accounting for cluster evolution than the reference SSP models (as shown in the top right-hand panel of Fig. 12), as $\Delta \log_{10}(M/L_V)$ shrinking from -0.24 to a range between -0.13 and -0.04 according to different assumptions. The differences among these three disruption-time-scale assumptions indicate that the clusters in our sample are in quite different evolutionary stages.

For instance, the four youngest clusters are more likely dynamically unevolved as given their comparatively long relaxation times relative to their current ages.

Clearly, if we combine mass underestimation discussed in Section 6.2.1, the offset in $\Delta \log_{10}(M/L_V)$ can become negligible, implicitly assuming the two dynamical effects are at least partly independent. The results shown in Fig. 13 allow us to conclude that internal/external dynamical effects plausibly account for the offset between our measured M/L_V and the SSP models shown in the upper panels of Fig. 12.

6.3 Cluster dissolution in the LMC and SMC

For star clusters in a tidal field, the total disruption time depends on the cluster’s initial mass as $t_{95 \text{ per cent}} \propto M_i^\gamma$, based on both N -body simulations (Baumgardt & Makino 2003) and observations (e.g. Boutloukos & Lamers 2003; Lamers et al. 2005b). The index γ has been found to be 0.62 and 0.60 ± 0.02 , respectively, from N -body simulations (Baumgardt & Makino 2003) and observations of solar-neighbourhood open clusters (Lamers, Gieles & Portegies Zwart 2005a). A scaling factor, t_4 , specifies the disruption time of a $10^4 M_\odot$ star cluster within its host galaxy.

Adopting a model in which both t_4 and the cluster formation rate (CFR) are constant, Boutloukos & Lamers (2003) derived $\log t_4 = 9.90 \pm 0.20$ and $\gamma = 0.61 \pm 0.08$ from the analysis of 314 SMC clusters located within 4 kpc from that galaxy’s centre. For the LMC, Parmentier & de Grijs (2008) used the same approach to constrain t_4 but with γ set to a fixed value of 0.62. They concluded only that $t_4 \geq 1$ Gyr, principally due to an apparent steady increase of the CFR in the LMC over the past 5 Gyr, negating one of the assumptions of the analysis.

We show here that we can use our M/L_V results to constrain t_4 in both galaxies. We start with the Anders et al. (2009) models as shown in the left-hand panel of Fig. 13. For a given age and metallicity, we can read off a M/L_V corresponding to a given disruption time. Using equation (2) from Anders et al. (2009), we can write for a given cluster

$$M'_{\text{tot}}(t) = f(t) \cdot M_i = f(t) \cdot 10^4 M_\odot \cdot \left(\frac{t_{95 \text{ per cent}}}{t_4} \right)^{1/\gamma} \cdot \left[\frac{\mu_{\text{ev}}(t_{95 \text{ per cent}})}{\mu_{\text{ev}}(t_4)} \right]^{-1}, \quad (4)$$

where t is the cluster’s age. The parameter μ_{ev} specifies the mass-loss of a cluster due to standard stellar evolutionary effects (e.g. mass-loss). The function $f(t)$ specifies the remaining bound mass fraction of a clusters when both stellar evolution and dynamical effects are

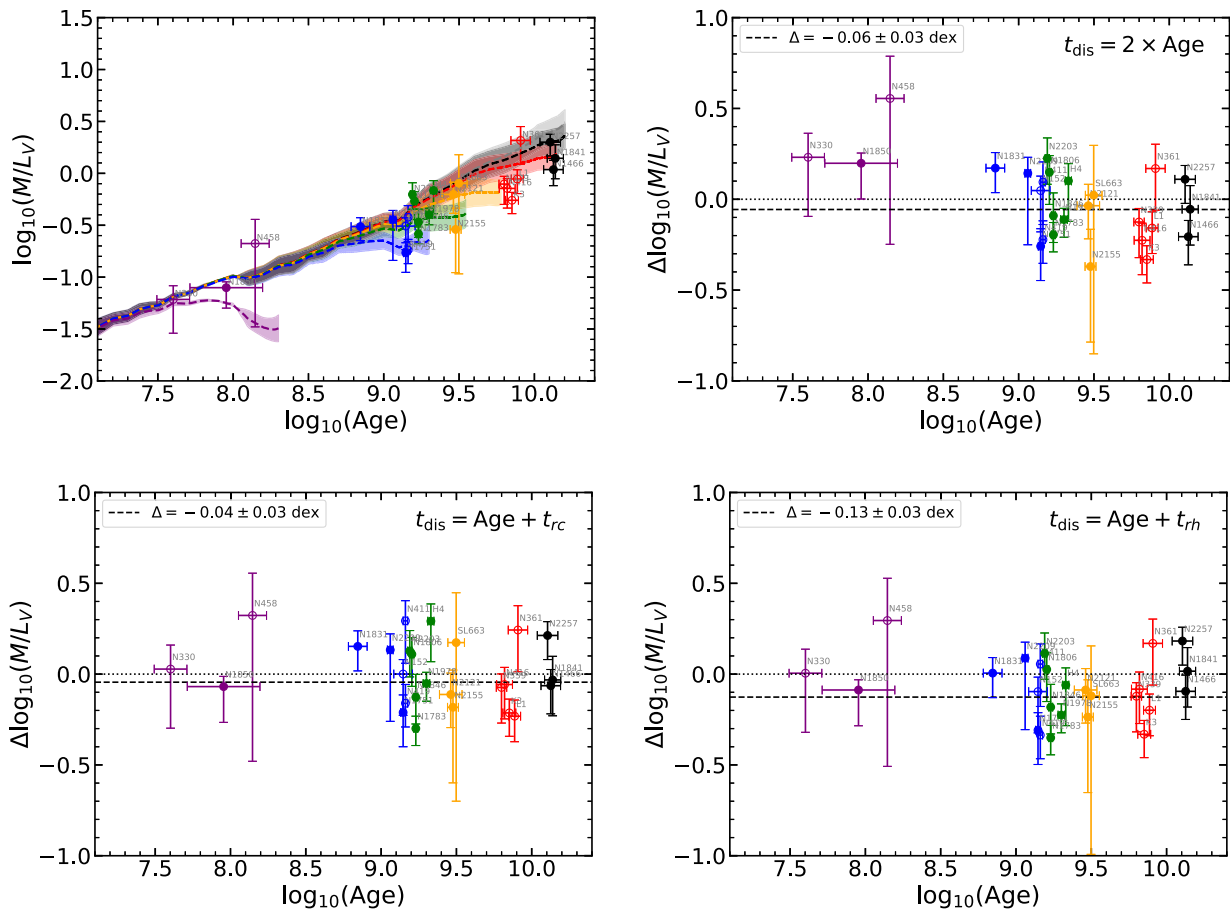


Figure 13. The same as the upper panels of Fig. 12, but now the coloured bands denote results from the synthesis models (Anders et al. 2009) that account for cluster dynamical evolution (see Section 6.2.2). In the top left-hand panel, the colour bands denote a range of evolutionary tracks over metallicity from -2.0 (lower band boundaries) to -0.7 (upper band boundaries) dex, and the coloured dashed curves denote the mean evolutionary tracks in this metallicity range. Different colours correspond to disruption times of twice the bin ages listed in Table 10. In the remaining panels, the $\Delta \log_{10}(M/L_V)$ values were calculated in a cluster-by-cluster manner using the age and metallicity of each cluster and compared with the Anders et al. (2009) models under the assumption that the total disruption time is double the age of the cluster (top right), the sum of the age of the cluster and its current core relaxation time (bottom left), and the sum of the age of the cluster and its current half-mass relaxation time (bottom right), respectively.

considered. Given a value for t_4 and γ , we can use this procedure to generate isochrones in the $(M/L_V)' - M'_{\text{tot}}$ plane.

The left-hand panel of Fig. 14 shows a set of such isochrones for $t_4 = 8.0$ Gyr (the measured value for the SMC by Boutloukos & Lamers 2003), while the right-hand panel shows results for $t_4 = 0.8$ Gyr (consistent with the lower limit value for the LMC clusters by Parmentier & de Grijs 2008). In both cases, we have adopted $\gamma = 0.62$ (Boutloukos & Lamers 2003). The isochrone colours correspond to the adopted mean ages and metallicities for the bins in which the clusters have been assigned (see the caption). Since the clusters of a given age bin vary in age and metallicity, we have shifted their positions in Fig. 14 (relative to the lower left-hand panel of Fig. 12) by using the reference SSP model (i.e. the model with a disruption time of 200 Gyr in Anders et al. 2009) to determine the small shifts in M and L_V associated with the shift in age and metallicity of each cluster to the corresponding bin values. The ‘prime’ notation (M'_{tot} and $(M/L_V)'$) is meant to emphasize that the plotted values have been adjusted from the results shown in Fig. 12.

In both panels of Fig. 14, we can see that $(M/L_V)'$ is constant for the highest mass and youngest clusters as these systems have not yet attained internal energy equipartition; thus their $(M/L_V)'$ values are nearly the same as expected for the reference SSP models (see

Fig. 12). For the $t_4 = 0.8$ Gyr models (right-hand panel), it can also be seen that as age increases, $(M/L_V)'$ can increase with decreasing mass. This reflects the fact that, at any given age, low-mass clusters will have lost more low-mass stars due to energy equipartition (Kruijssen 2008). For the lowest mass clusters (e.g. the black and red dotted lines in the right-hand panel of Fig. 14), the increasing fraction of bound stellar remnants near the end stages of cluster dissolution (Anders et al. 2009) causes a rapid increase in $(M/L_V)'$. This also implies that clusters found near the minima of the isochrones are very close to complete dissolution.

Note that for the larger value of t_4 (8.0 Gyr; left-hand panel in Fig. 14), the isochrones tend to run above the data for clusters of corresponding age. For the smaller t_4 value (0.8 Gyr; right-hand panel), the isochrones systematically match the cluster data better in both the SMC and LMC. The reduced χ^2 values are 5.15 and 2.16 for the cases of $t_4 = 8.0$ Gyr and 0.8 Gyr, respectively, with a degree of freedom of 25. This indicates that most clusters in our sample have evolved in a relatively strong tidal field with a small t_4 value. However, there are two possible exceptions to this conclusion. First, the oldest clusters (in black) appear to agree better with the larger t_4 value. These clusters, all associated with the LMC, are located furthest from the galaxy centre. This suggests that they may

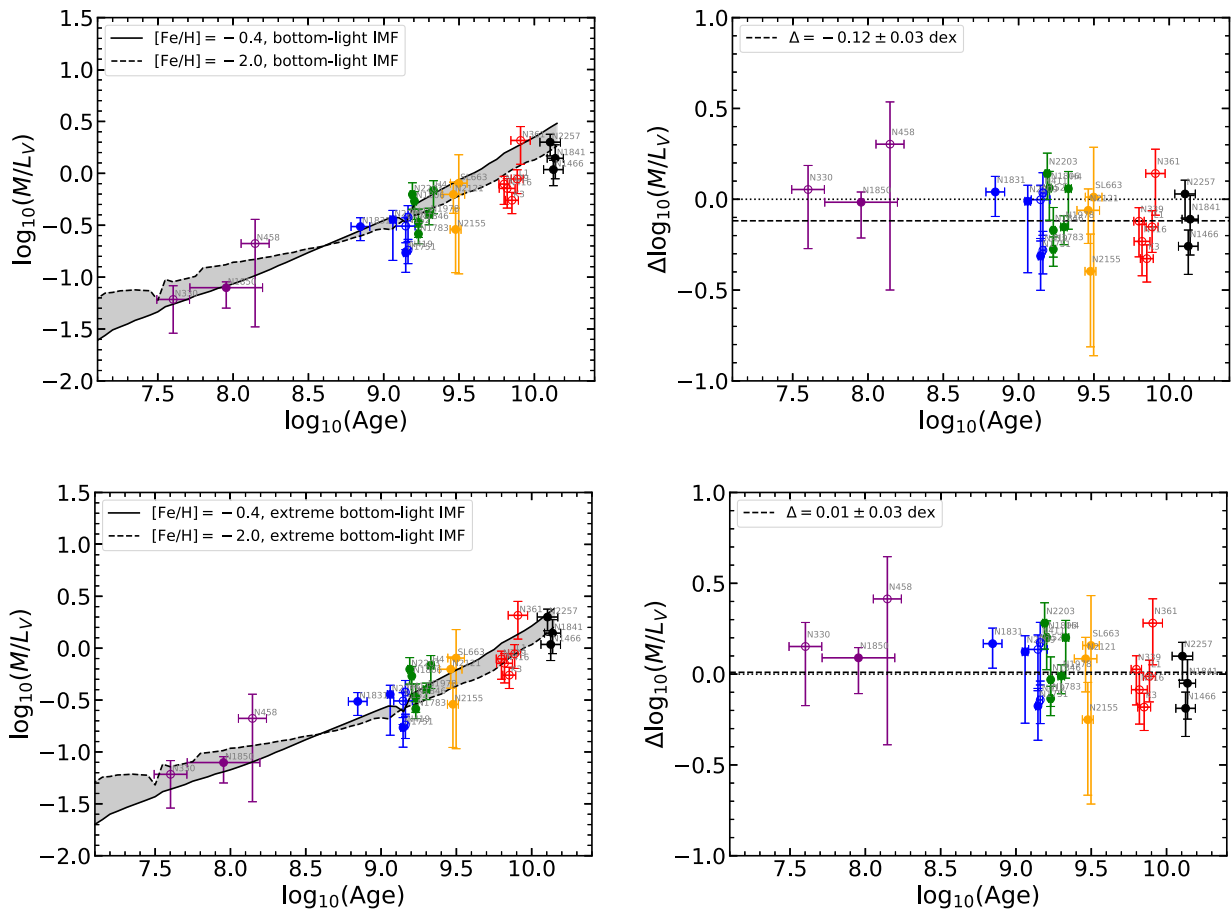


Figure 15. The same as the upper panels of Fig. 12, but now the solid and dashed curves and their enclosed grey band denote the evolutionary tracks of SSP models with a bottom-light (top panels) and an extreme bottom-light IMFs (bottom panels), respectively. These two IMFs are defined in Section 6.4. The $\Delta \log_{10}(M/L_V)$ values in the right-hand panels were again calculated in a cluster-by-cluster manner using the age and metallicity of each cluster similar to those shown in Fig. 12. The colours of the points correspond to the age bins described in Table 10.

these integrated-light measurements as well as M/L_V estimates of MC clusters based on individual-star spectroscopy of intrinsic precision similar to that of our study (Fischer et al. 1992a, b, 1993; Suntzeff et al. 1992; McLaughlin & van der Marel 2005; Mackey et al. 2013; Kamann et al. 2018b; Patrick et al. 2020). This comparison highlights some key points. First, all individual-star results seem to agree systematically over the full range of ages explored by the data. Moreover, these data run roughly parallel to the model expectations. Secondly, the integrated-light measurements appear to define a relation that is considerably shallower than the SSP models or the individual-star measurements. Thirdly, modified SSP models based on a bottom-heavy IMF ($\alpha_1 = \alpha_2 = 2.8$ and $\alpha_3 = 2.3$, plotted in Fig. 16) agree well with the Zaritsky et al. (2012, 2014) integrated-light results for clusters younger than about 3 Gyr. For older clusters, the integrated-light and individual-star M/L_V results broadly agree (see Section 5.3).

Integrated-light studies tend to favour high-concentration clusters with moderate to bright central surface brightnesses, while individual-star studies tend to employ more diffuse and larger clusters where obtaining spectra of distinct stars is more feasible. Moreover, the integrated-light spectra tend to only consist of contributions from the innermost regions of the clusters. These may point to a possible physical distinction between the clusters that reveals real IMF variations. However, the systematic tendency of the integrated-light results to run high compared to individual-star results, particularly

for clusters younger than about 3 Gyr, suggests a more prosaic explanation. For instance, systematic effects in the integrated-light studies have the potential to inflate the line widths (such as blends and focus errors), leading to overestimates of the velocity dispersion (Zaritsky et al. 2012). Clearly, more studies of clusters observed using both techniques would help explore the nature of this apparent discrepancy.

7 SUMMARY AND CONCLUSIONS

In this paper, we have presented *Magellan*/M2FS observations of (mostly) red giants in and around 26 Magellanic Cloud star clusters (10 in the SMC and 16 in the LMC) chosen to span the range from ~ 100 Myr to ~ 13 Gyr in age, and from -2.0 to -0.4 in $[\text{Fe}/\text{H}]$. We employed an improved version of the data reduction process (described in S19) to extract from the raw data 3137 stellar spectra of 2901 distinct targets. Using stellar effective temperatures estimated with *Gaia* DR2 G -band magnitudes (see Section 3.2), we applied a Bayesian methodology to obtain radial velocities, metallicities, and surface gravities from these spectra (see Section 3.1 and Song et al. 2017). These parameters were used to identify peculiar targets, such as C stars, binary/blended stars, extreme M supergiants, etc. (see Section 3.4). Combined with previously published velocities of individual stars in the clusters of our sample, we have produced a

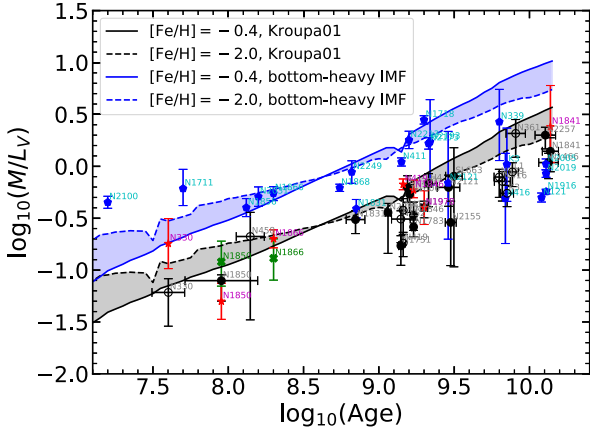


Figure 16. The same as the top left-hand panel of Fig. 12, but with all M/L_V values from this work shown in black. Overplotted are the M/L_V results from previously published studies using either integrated-light spectroscopy (in blue; Zaritsky et al. 2012, 2014) or individual stellar spectra (in red and green). The green data points denote the re-calculated results of McLaughlin & van der Marel (2005) by fitting K66 models to previously published velocity dispersion results (i.e. NGC 1866 from Fischer et al. 1992a and NGC 1850 from Fischer et al. 1993), and hence are labelled differently from other individual-star studies shown in red (Fischer et al. 1992b; Suntzeff et al. 1992; Mackey et al. 2013; Kamann et al. 2018b; Patrick et al. 2020). We also plot the evolutionary tracks of SSP models with a bottom-heavy IMF (blue curves and band) to illustrate the trend of integrated-light results for clusters younger than about 3 Gyr (blue pentagons).

kinematic data set of 2787 stars suitable for dynamical and chemical analyses.

Using this sample, we have determined membership probabilities of individual stars in each cluster using an EM algorithm (Walker et al. 2015a, b; Song et al. 2019) with the assumption that cluster members are spatially and kinematically distributed as expected for a single-mass K66 model (King 1966). The EM algorithm we used assumes that superimposed on the cluster is a spatially uniform field population that follows a kinematically much broader Gaussian distribution than the cluster population. In order to properly account for the influence of likely non-members, we followed the same approaches developed in S19 to assign cluster membership probabilities for the stars in each cluster sample (Section 4). We found that for five clusters, the so-called PM50 samples (comprised of stars with membership probabilities greater than 50 per cent) still have potential field interlopers. We developed a related approach, PM50', to identify cases where significant contamination by a single unflagged non-member appears to be present (see Section 5.1). Using the cluster members in the resulting PM50 or, for a few clusters, PM50' samples, we obtained projected central velocity dispersion of each cluster. From these, we have derived total masses, M/L_V ratios and mean metallicities for all 26 clusters in our sample.

Our results exhibit readily understandable trends of M/L_V with cluster age, metallicity, and mass (see Section 6.1). When compared with the reference SSP models, we found that our empirical M/L_V values are about 40 per cent (-0.24 dex in $\Delta \log_{10}(M/L_V)$) lower than model predictions over the full range of ages exhibited by the clusters in our sample. We explored the origin of this offset by considering two specific dynamical effects (Section 6.2). First, the single-mass K66 model we adopted do not account for energy equipartition and, hence, mass segregation within clusters. Consequently, these models will tend to underestimate a cluster's total mass compared to more realistic multimass models. N -body simulations of star clusters (Sollima et al.

2015; Hénault-Brunet et al. 2019) suggest that this only partially accounts for the offset as we found (about 0.08 dex). Secondly, using modified SSP models that account for cluster evolution in a tidal field (Anders et al. 2009), the M/L_V offset is significantly reduced for reasonable cluster disruption time-scale assumptions.

The observed masses and M/L_V ratios of the clusters in our sample were used to try to constrain time-scales of cluster dissolution in the MCs using a simple tidal disruption model (see Section 6.3; also Baumgardt & Makino 2003; Boutloukos & Lamers 2003; Parmentier & de Grijs 2008). Our results suggest that external tidal dynamical effects lead to relatively rapid dissolution in both galaxies ($t_4 \sim 0.8$ Gyr, where t_4 is time to disrupt half of a sample of clusters with initial mass $10^4 M_\odot$). One exception is, perhaps, clusters located well outside the main body of the LMC where a longer t_4 is indicated. For the LMC, this model is consistent with the assumption of a non-constant cluster formation rate (CFR) (Parmentier & de Grijs 2008). Our analysis suggests that the LMC cluster NGC 2155 – with an age of 3.0 Gyr – may be close to total dissolution, having already lost 65–90 per cent of its initial mass. More detailed kinematic studies of this cluster and other LMC clusters near the temporally more recent edge of the ‘age gap’ (Bertelli et al. 1992; Girardi et al. 1995; Olszewski et al. 1996) may provide important constraints on the CFR of this galaxy. For the SMC, our analysis suggests a short mean cluster disruption time-scale (possibly non-constant CFR in the past), somewhat in contrast to the conclusions of e.g. Boutloukos & Lamers (2003) who assumed a constant CFR and comparatively long cluster disruption time-scale.

We also considered the effect of varying stellar IMFs among star clusters to explain the offset between observed and model M/L_V values (Section 6.4) and found that an extreme bottom-light IMF could, by itself, almost fully account for the offset (Fig. 15). In contrast, the fact that our observed M/L_V ratios run lower than SSP models with a Kroupa (2001) IMF strongly disfavours a bottom-heavy IMF for these clusters.

A key conclusion of this study is that both dynamical effects and IMF variations can account for the M/L_V values we have measured for our cluster sample (see Figs 13, 14, and 15). It is worth noting, however, that while dynamical effects must be affecting the evolution of MC star clusters, it is not as clear that IMF variations can or must be present. Thus, we favour the dynamical modifications to SSP models as described in Sections 6.2.2 and 6.3 as the more likely reason to account that our M/L_V results run systematically below the predictions from the reference SSP models. To the extent that IMF variations may be present, our results require these to be in the form of a bottom-light mass function since bottom-heavy IMFs would cause M/L_V values to run higher than the reference models. As long as dynamical or minor IMF variations are allowed, we find that present-day SSP models such as those used for our analysis do a remarkably good job of explaining the integrated M/L_V values we observe for MC clusters.

We end by noting that our conclusion that SSP models perform well is tempered by a few factors. First, the dynamical models in our analysis are fairly simple. More realistic models could take into account a spread of stellar masses, anisotropy, and possible rotation within clusters. This increased sophistication would come at the price of additional parameters that are likely only weakly constrained by our cluster data sets. Secondly, we have not attempted to account for binary stars in our analysis. We have noted the presence of a some binaries of moderate velocity amplitude in two clusters in our sample (see Section 4.3). Such systems are ‘easy’ cases since they are most likely to be flagged as non-members. More problematic are binaries with small velocity amplitudes comparable to the clusters’ internal

dispersions of a few km s^{-1} (see Spencer et al. 2018). Producing empirical constraints on the frequency of such binaries would be a formidable task. One can imagine that missed binaries would systematically inflate the observed cluster velocity dispersions – and hence their *M/L* ratios – making the offset between models and data larger than 40 per cent shown in Fig. 12. But the magnitude of the effect remains poorly constrained given the large uncertainties in the relevant binary-population parameters (see Spencer et al. 2018).

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DATA AVAILABILITY

The full version of Table 5 in this article is in its online supplementary material.

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SUPPORTING INFORMATION

Supplementary data are available at *MNRAS* online.

Table S1. Sample of 3095 Targets from 26 Star Clusters.

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APPENDIX A: DETERMINING CLUSTER CENTRES

As noted in Section 2.1, accurate cluster centres are crucial for both photometric and dynamical analysis. Carvalho et al. (2008) showed that errors in the cluster centre can cause significant offsets in the resulting structural parameters, especially the central surface brightness, measured from the surface brightness profile. We explore this further at the end of this section to determine specifically how sensitive our final masses are to the centring errors in the cluster centres. As we shall show, the net impact of centring errors described below appears to be generally small, mostly negligible given other sources of errors for the derived dynamical properties of the clusters in our sample.

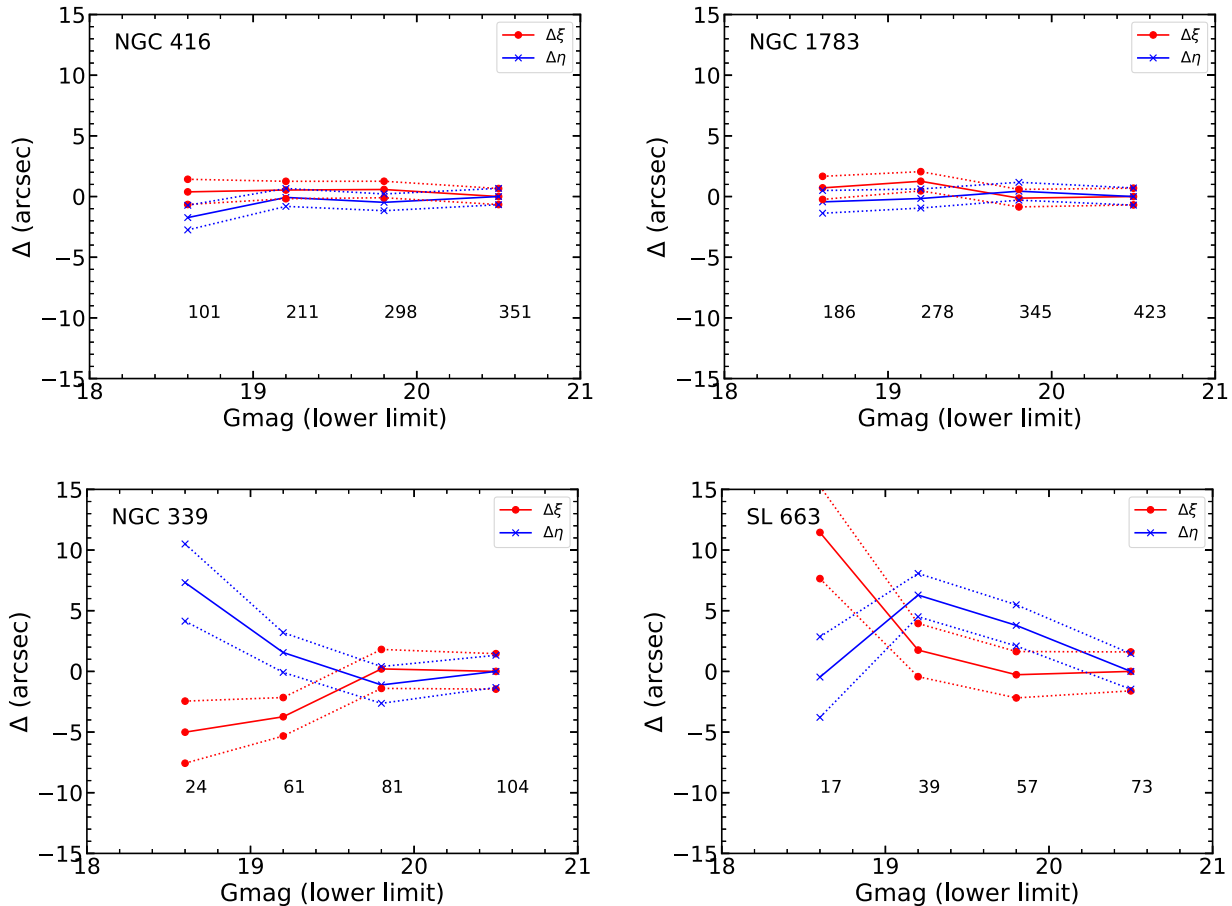


Figure A1. Residuals in RA ($\Delta\xi$, red dots) and DEC ($\Delta\eta$, blue dots) as a function of *Gaia* DR2 limiting magnitude for targets in the MC clusters of this study. Four representative clusters – NGC 416, NGC 1783, NGC 339, and SL 663 – are presented to illustrate the stability and precision of the cluster centres determined as described in Appendix A. In all panels, the solid red and blue lines represents the best-fitting value of RA (ξ) or DEC (η), respectively, while the dotted lines show the corresponding $1\text{-}\sigma$ error ranges. The numbers below each set of symbols in each panel give the number stars within 30 arcsec in radius from the cluster centre to that magnitude limit. The upper two clusters have comparatively large samples of stars in the *Gaia* DR2 and their centre positions remain fairly stable as fainter targets are introduced. The lower panels are for sparser clusters for which the centre positions vary more strongly with magnitude. These plots suggest that cluster centres determined in this manner become stable when samples of 70 or more stars are used.

Many previous studies have determined cluster centres from photometric data as part of their analysis of the surface brightness profiles (e.g. Carvalho et al. 2008; Glatt et al. 2009; Mackey et al. 2013). However, data on cluster centres is not available in the literature for all clusters in our sample. To obtain this information a more convenient approach is to take cluster centres from online astronomical data bases (e.g. Mackey & Gilmore 2003a,b), such as SIMBAD (Wenger et al. 2000). A major limitation of this approach, however, is that the methodology, accuracy, and uncertainties of the centres cited in this approach are generally not provided nor are they assured of being internally consistent. Some centre positions listed for NGC 1850 and NGC 2209, for example, appear to be significantly offset from their locations on DSS images. For this reason, we chose to re-determine the centres of all the clusters in our sample using the *Gaia* DR2, as described next.

For a given cluster, we first selected all stars within 7 arcmin from the centre coordinate listed on SIMBAD, brighter than 20.5 mag in *G*-band and with parallax less than 0.2 mas. This selection radius is considerably larger than the core (and often tidal) radii of our clusters. The selected stars were then used to compute an initial estimate of the position of the cluster centre based on the mean positions in RA and Dec. The position was improved over four

iterations (in a process similar to the description in Glatt et al. 2009) as the selection radius was decreased. For example, a second position estimate was calculated using all stars within a radius of 4 arcmin from the initial guess position. In subsequent iterations, the selection radii were halved from the previous iteration; the last iteration used a selection radius of 0.5 arcmin (for reference, most core radii of the clusters in our sample are smaller than this value; see Table 2). The primary reason for simply calculating the mean positions is that the *Gaia* DR2 remains impressively complete near the centres of the target clusters and hence there is a strong positional weighting inherent in the stellar samples used to determine the cluster centres.

We adopted the coordinates of the last iteration for each cluster. The resulting uncertainties of the cluster centres using this procedure were estimated using the standard deviation of RA or Dec of all stars in the last iteration divided by the square root of the sample size. These uncertainties are typically between 0.6 and 2.0 arcsec in both RA and Dec. The adopted cluster centres are listed in columns 3 and 4 of Table 2.

To address any bias or funneling effect in our centre determination, we re-derived three extra sets of cluster centres with the different lower magnitude limit of stars – instead of 20.5 mag, we also used 18.6, 19.2, and 19.8 mag. Fig. A1 provides some examples of the

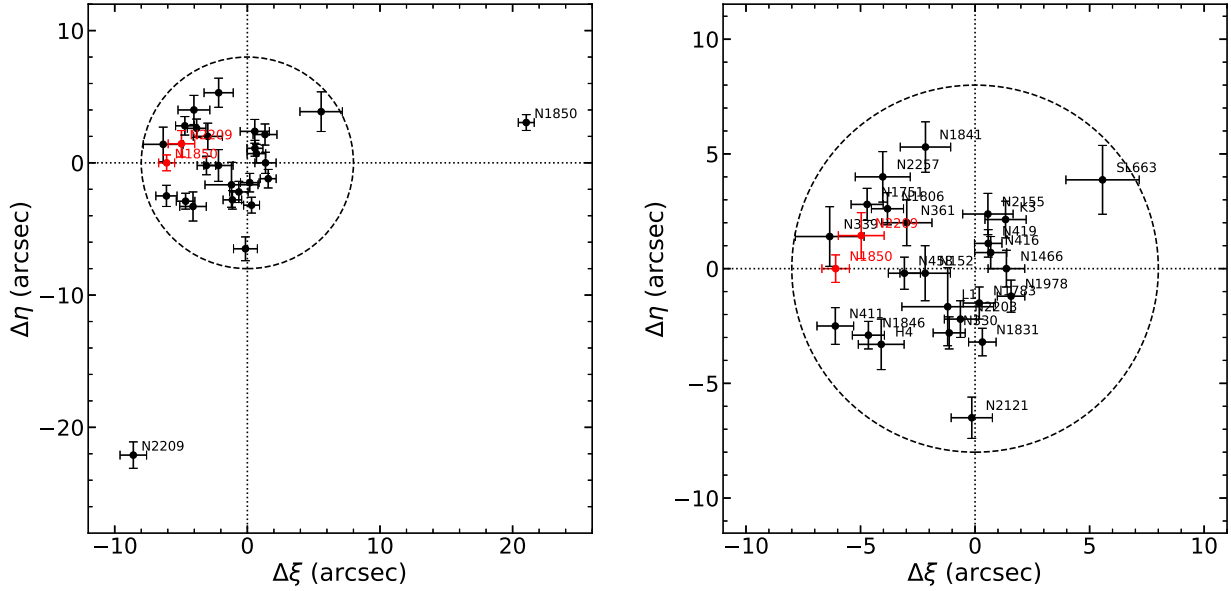


Figure A2. A comparison in the $(\Delta\xi, \Delta\eta)$ plane of the cluster centres determined using the *Gaia* DR2 catalogue (Appendix A) with centres provided by SIMBAD (Wenger et al. 2000, in black), where Δ is defined as the SIMBAD minus the *Gaia* DR2 centre positions. The red dots correspond to our centres compared with centre coordinates from other studies (Milone et al. 2018 for NGC 1850 and Mucciarelli et al. 2006 for NGC 2209). The circles in both panels are 8 arcsec in radius. The left-hand panel shows all the clusters studied in this work, including the two most deviant cases, NGC 1850 and NGC 2209. The right-hand panel highlights the distribution of central position residuals for all clusters in our sample. Further details are provided in Appendix A.

residuals in RA (red dots), and DEC (blue crosses) from the centres determined using 20.5 mag as the lower magnitude limit. For most clusters (e.g. the top two panels of Fig. A1), the standard deviations of the four best-fitting values in either RA or DEC are below 3 arcsec, indicating good stability. For a few clusters, i.e. Kron 3, Lindsay 1, NGC 339, NGC 361, NGC 411, and NGC 1850, at least one of their standard deviations in RA and Dec falls between 3 and 5 arcsec, which is mostly due to the small number of stars brighter than 18.6 mag (see e.g. the bottom left-hand panel of Fig. A1). The worst case is SL 663, for which the values are 6.7 and 4.3 arcsec in RA and DEC, respectively. As shown in the bottom right-hand panel of Fig. A1, the results can also be attributed to the small number of stars brighter than 18.6 mag.

Fig. A2 compares the derived cluster centres with those provided by SIMBAD in the $(\Delta\xi, \Delta\eta)$ plane. The $\Delta\xi$ and $\Delta\eta$ values represent the positional offsets in RA and DEC, respectively, from the SIMBAD centres to the *Gaia* DR2 centres. For NGC 1850 and NGC 2209, the new centres are offset more than 20 arcsec from the old ones (see the left-hand panel of Fig. A2), but they are consistent with the central locations on the DSS images. Except for these two outliers, the remaining cluster centre offsets are all within 8 arcsec (see the right-hand panel of Fig. A2), which is typically multiple times smaller than the King radii listed in Table 2, column 8. We have also compared the new centres of NGC 1850 and NGC 2209 with those from other sources not cited by SIMBAD (Mucciarelli et al. 2006; Milone et al. 2018), and found that they agreed to the same level of precision as the other clusters in our sample (see the red dots in Fig. A2).

To further test the influence of cluster centre offsets on the central velocity dispersion, we applied the dynamical analysis described in Section 5 to the samples of NGC 419 and NGC 1846 published in S19. Using 100 random positions that are 10 arcsec from the original cluster centre, the derived central velocity dispersions are negligibly different from the reported dispersions reported in S19. For this study, we adopt the cluster centres derived from the *Gaia* DR2 data.

APPENDIX B: BACKGROUND SUBTRACTION

For Method A, we model the background in the vicinity of a cluster as

$$n(r) = n_k \left[\frac{1}{\sqrt{1 + (r/r_0)^2}} - \frac{1}{\sqrt{1 + (r_i/r_0)^2}} \right]^2 + n_b, \quad (\text{B1})$$

where $n(r)$ is the median counts obtained from all background sources that contribute to our spectra, n_k is the peak counts that come from a K62 profile with structural parameters equal to those of the cluster being analysed (see Table 2), and n_b is the counts the extended background from non-cluster sources (for instance, unresolved light from LMC/SMC field stars, telluric emission, moonlight). If spectra from dedicated sky fibres and offset observations are both available, we can solve for n_k and n_b given the cluster structural parameters. The key for this method is to have enough sky positions near the cluster centre – typically from the offset observations – to effectively constrain n_k .

Strictly speaking, n_k and n_b in equation (B1) should be determined for each wavelength to account for spectral variations of the background with distance from a cluster’s centre. In practice, our background sampling is too sparse in the inner regions of nearly all of our target clusters to attempt this. We therefore took the spectral shape of the background spectrum to be equal to the mean of the spectra from all the dedicated sky fibres (that is, $(n_k/n_b)_\lambda = \text{Constant}$ for all wavelengths). Clusters for which Method A was applied are flagged with an ‘A’ in Table 3.

When no offset observations, are available, determining n_k and n_b is less precise due to poor or non-existent sampling of the background near the a cluster’s core. For these cases we used ‘Method B’ (S19) in which n_k is estimated for a given ‘Target’ cluster by assuming that the ratio of this parameter for the cluster divided by the value of the parameter in a ‘Reference’ cluster is the same as the ratio of the photometric central surface brightnesses, $\Sigma_{V,0}$, of the clusters

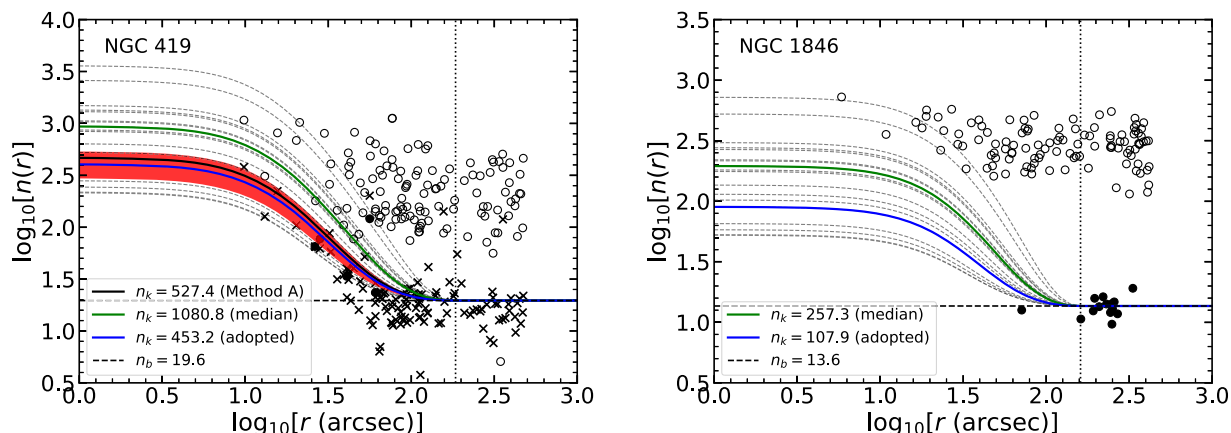


Figure B1. Two examples of background subtraction. In both panels, symbols represent the median total counts of target spectra (open circles) and dedicated sky fibres (filled circles). The crosses in the left-hand panel represent target fibres when the telescope was offset as described in Section 2.4.1. For NGC 419 (left-hand panel), the black curve shows the best-fitting profile determined using equation (B1) (defined as ‘Method A’ in Section 2.4.2) which employs both the dedicated sky fibres and offset-sky observations for that cluster. The y-intercepts of the horizontal dashed line and the solid black curve denote the values of the background and central mean counts (n_b and n_k , respectively; see equation (B1)) of the fitted profile. The red band shows the central 68-percentile (taken as $\pm 1\sigma$) determined from a bootstrap error estimator that takes into account observational scatter on the estimated value of n_k . The right-hand panel illustrates the application of equation (B2) (Method B, Section 2.4.2) to our data for NGC 1846. Here, the background mean counts, n_b , are determined from the dedicated sky fibres (filled circles) located outside the cluster tidal radius, r_t (denoted as a vertical dotted line for both clusters). The central background profile counts, n_k , are obtained from pairwise applications of equation (B2) with the 17 clusters in the sample for which Method A was applicable (see Table 3). The dashed grey profiles each represent one application of equation (B2) to a specific cluster. The green solid line follows the median (central) curve. The left-hand panel also shows Method B results applied to NGC 419 where we use the outermost offset-sky fibres to estimate n_b and then apply equation (B2) as for NGC 1846. Note that the median profile (green line) runs above the black curve obtained using Method A. When this same procedure is applied to all Method-A clusters, we find the best average agreement between Method A and Method B profiles to correspond to the ‘minus 1- σ ’ profile (the fifth-lowest profile of the 17 plotted, shown as a blue curve in both profiles).

adjusted by exposure time. That is,

$$\frac{n_{k, \text{Target}}}{n_{k, \text{Reference}}} = \frac{\sum_{V, 0, \text{Target}}}{\sum_{V, 0, \text{Reference}}} \cdot \frac{t_{\text{exp, Target}}}{t_{\text{exp, Reference}}}. \quad (\text{B2})$$

In practice, a Reference cluster is one for which n_k was determined using Method A; if we have N Reference clusters, we can produce a distribution of estimates for $n_{k, \text{Target}}$.

Method B is somewhat crude in that it ignores possible transparency variations between observations of different clusters and it relies on the precision of the central surface brightnesses of the clusters in the sample, typically about 10–50 per cent. But the method has the advantage of being applicable to all the clusters in the sample. A corollary benefit is that we can use any clusters that are suitable for background determination using equation (B1) (Method A) and apply equation (B2) (Method B) as a check on how well the latter method works.

Fig. B1 shows the application of equation (B1) (Method A) to NGC 419, one of the clusters we analysed in S19. In this case, n_b and n_k can be determined reliably from dedicated sky fibres and offset observations; the resulting background profile is shown as a black solid curve in the left-hand panel of Fig. B1. The right-hand panel of this figure illustrates how we apply Method B (equation B2) for the case of NGC 1846, the other cluster in S19. Here, n_b is well-determined from the dedicated sky fibres (filled circles in Fig. B1) in the field surrounding the cluster, while n_k is poorly constrained due to the lack of offset observations to sample the background in the inner parts of the cluster. In this instance, n_k has been estimated by applying Method B (equation B2) in a pairwise manner to all 17 clusters for which Method A could be applied. These background profiles are shown as dashed grey solid lines. These profiles all have the same value of n_b as derived from

the NGC 1846 dedicated sky fibres located far from the cluster centre, but all have distinct values of n_k values derived by applying equation (B2).

As previously noted, we can apply Method B to clusters for which Method A is also applicable. In the case of NGC 419 (left-hand panel, Fig. B1), the median background profile of the set of profiles obtained using Method B (shown as a green solid curve) runs significantly above the one profile obtained using Method A (black line). Applying this test to all the clusters in our sample for which Method A was applicable, we found that the background profile (shown as the blue solid curve in Fig. B1) was consistently located below the median profile.

APPENDIX C: REJECTED AND/OR UNUSUAL M2FS SPECTRA

C1 Examples of M2FS spectra rejected by the SK cut

As described in Section 3.4.1, all but three of the stars rejected for skew/kurtosis have low average S/N of 1.2 ± 0.1 . Among these remaining 68 stars, three appear to have spectra consistent with those of carbon (C) stars (more on these in Appendix C2) at low average S/N (~ 1.2 for these three; see Fig. C1). The additional three spectra flagged in the SK cut but that have fairly high mean S/N (25 ± 8) either contain very few identifiable spectral features or possibly have broadened or blended features. Not surprisingly, the Bayesian analysis was unable to settle on single well-defined LOS velocities in these cases, resulting in large skew/kurtosis indices. For reference, the spectra of these three higher S/N objects that failed the ‘SK cut’ are also shown in Fig. C1.

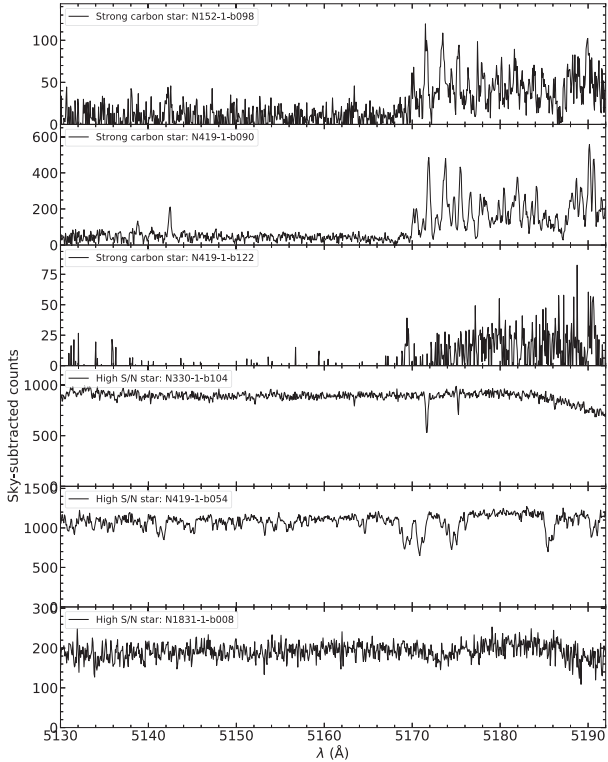


Figure C1. Examples of M2FS spectra of stars rejected by the SK cut described in Section 3.4.1. The three top panels appear to be strong C stars at low S/N (more on these in Section 3.4.2). The bottom three panels show stars that have high mean S/N values, but were still rejected by the SK cut due to having too few or blended lines.

C2 Examples of M2FS spectra with anomalous velocity uncertainties

As discussed in Section 3.4.2, we inspected each of the 62 stars with anomalous velocity uncertainties and found that we could classify them into the following distinct categories:

C2.1 Carbon stars

A total of 21 C stars were identified based on the presence of a clear C_2 Swan band feature at $\lambda 5165 \text{ \AA}$ (Johnson 1927; King 1948). Among these stars, 17 have quite strong absorption features (the first spectrum in Fig. C3 shows a typical example), while the remaining four C stars have relatively weak Swan-band absorption (the second spectrum in Fig. C3 shows a typical case). Since the spectral library we are using (Lee et al. 2008a,b) does not model C stars, it is unsurprising that our Bayesian analysis returned LOS velocities with large uncertainties relative to other stars at similar S/N.

C2.2 Hot stars

A total of eight stars appear to be hot stars that were forced to an incorrect low temperature using the method to fix T_{eff} described in Section 3.2. We chose not to single out these stars and allow their T_{eff} to float as this would be inconsistent with the strategy described in Section 3.2. Two examples, one also exhibiting moderate rotation, are shown in Fig. C3.

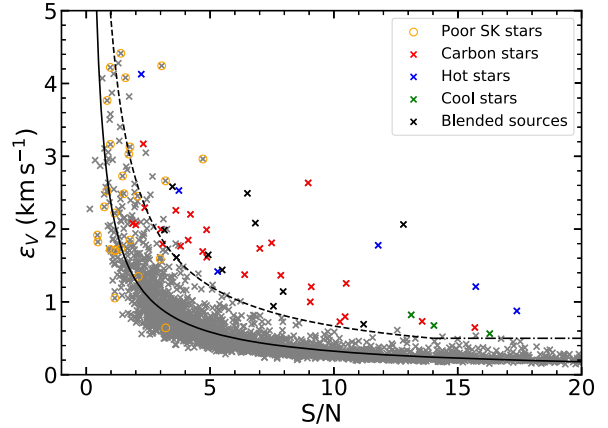


Figure C2. A similar plot as Fig. 5, but here in linear space and with anomalous stars highlighted by coloured symbols. The meanings of these symbols are given in the legend. Different symbols correspond to different types of anomalous stars identified in Section 3.4.2. The cases denoted as ‘poor SK stars’ are discussed in Section 3.4.1.

C2.3 Cool stars

A total of three stars appear to be cool giant stars with the feature of TiO bands above $\sim 5180 \text{ \AA}$ (the fifth spectrum in Fig. C3 shows a typical example). Similar to hot stars, cool stars were forced to an incorrect temperature – in this case too high – as described in Section 3.2. Moreover, the spectral library we used does not contain stars sufficiently cool that exhibit TiO bands. For these reasons, it is likely unsurprising that the Bayesian analysis returned poor LOS velocity uncertainties for these cases.

C2.4 Blended sources

A total of 12 stars are confirmed as sources with the features of more than one star (see the sixth spectrum in Fig. C3). The spectra in this category provide poor LOS velocities for a given S/N. We did not carry out dual-star fits as in S19 because these stars represent a fairly small sample that may have a different error distribution from the single-star fitted spectra.

C2.5 Statistical tail

A total of 18 stars appear to represent the tail of the distribution of the velocity errors relative to the fit line/power-law evident in Fig. C2. These stars represent about 0.6 per cent of the total sample of spectra, roughly consistent with what a 3σ cut would achieve for a normal distribution.

Due to the relatively high frequency of C stars in our sample of rejected spectra, we developed a more objective means of identifying candidate C stars regardless of where they lie in Fig. C2. Specifically, we measured the ratio of the fluxes in all the spectra of our sample on the red and blue sides of the C_2 band head (the first two spectra in Fig. C3 illustrate where these bands are located). This ratio cleanly identified all C stars in the velocity-error rejection region of Fig. C2, as well as seven apparent C stars below the rejection boundary. Four of these are strong C stars similar to the top spectrum in Fig. C3, and three are weaker C stars with spectra similar to the second panel of this figure. For consistency, we have removed these four strong C-stars

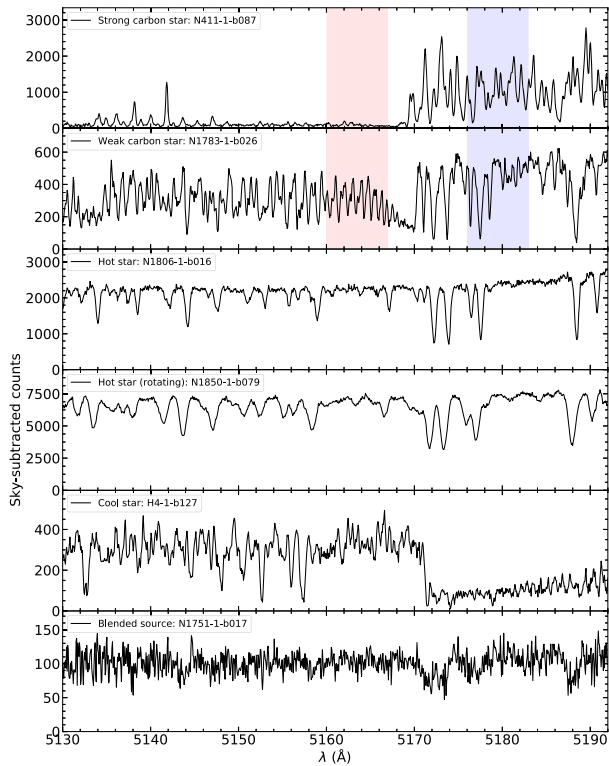


Figure C3. Examples of M2FS spectra of stars with anomalous velocity uncertainties as identified in Fig. C2. From top to bottom, we show the spectra of a strong C star, a weak C star, a star with high surface temperature (hot star), a similarly hot star but with moderate rotation ($v \sin i \sim 30 \text{ km s}^{-1}$), an M-type star with low surface temperature (cool star), and a blended source. For the two C stars, we indicate the two bands used to identify such stars as outlined in Appendix C2. A total of 62 targets with anomalous velocity uncertainties (as described in Section 3.4.1 and identified in Fig. C2) are excluded from our cluster dynamical analysis.

from our sample but have retained the three weak C-stars since they show clean atomic features similar to non-C stars of similar colour.

This method of identifying C stars found numerous even weaker C stars. We have chosen to accept these in the present sample as they are mildly affected by the C_2 features and in all cases their spectra appear to provide good LOS velocity estimates. We plan to discuss this expanded C-star sample in more detail in a future paper.

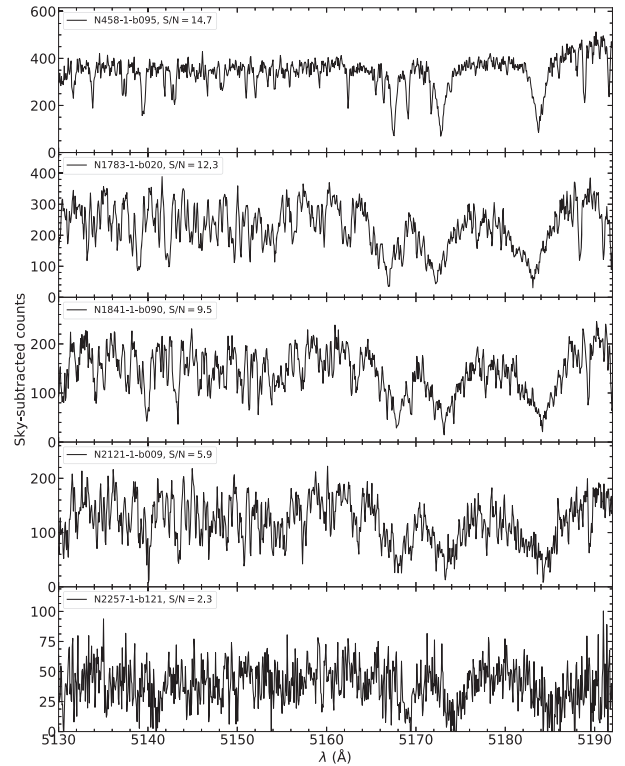


Figure C4. Example M2FS spectra of dwarf stars confirmed in our sample. From top to bottom, we show dwarf stars spanning a similar S/N range as those giants shown in Fig. 2; these dwarf stars and giants also have comparable colours. As discussed in Section 3.4.3 and Fig. 6, 84 dwarf stars with spectra similar to these were removed from the sample used for the dynamical analyses the clusters in this study.

C3 Examples of M2FS spectra from dwarf stars confirmed in our sample

In Section 3.4.3, we have removed 84 stars with surface gravity parameters $\log g \geq 3.2$. The distinctive spectral signatures of these stars are evident by comparing the dwarf-star spectra in Fig. C4 with spectra of giants of similar colours and S/N ratios shown in Fig. 2. As expected, the proper motions of these stars, as provided in *Gaia* DR2, exhibit a large spread that is offset from the mean proper motions of the MCs.

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